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Investing Cosmic rays with gamma-ray astronomy

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Outline

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1. Galactic Cosmic rays: direct measurement

2. Spatial distribution of Galactic Cosmic rays

3. Low energy Cosmic rays

4. PeV CR sources (PeVatrons)

Galactic Cosmic Rays (GCR)

Cosmic Rays: Relativistic particles (mainly protons) in interstellar medium (ISM)

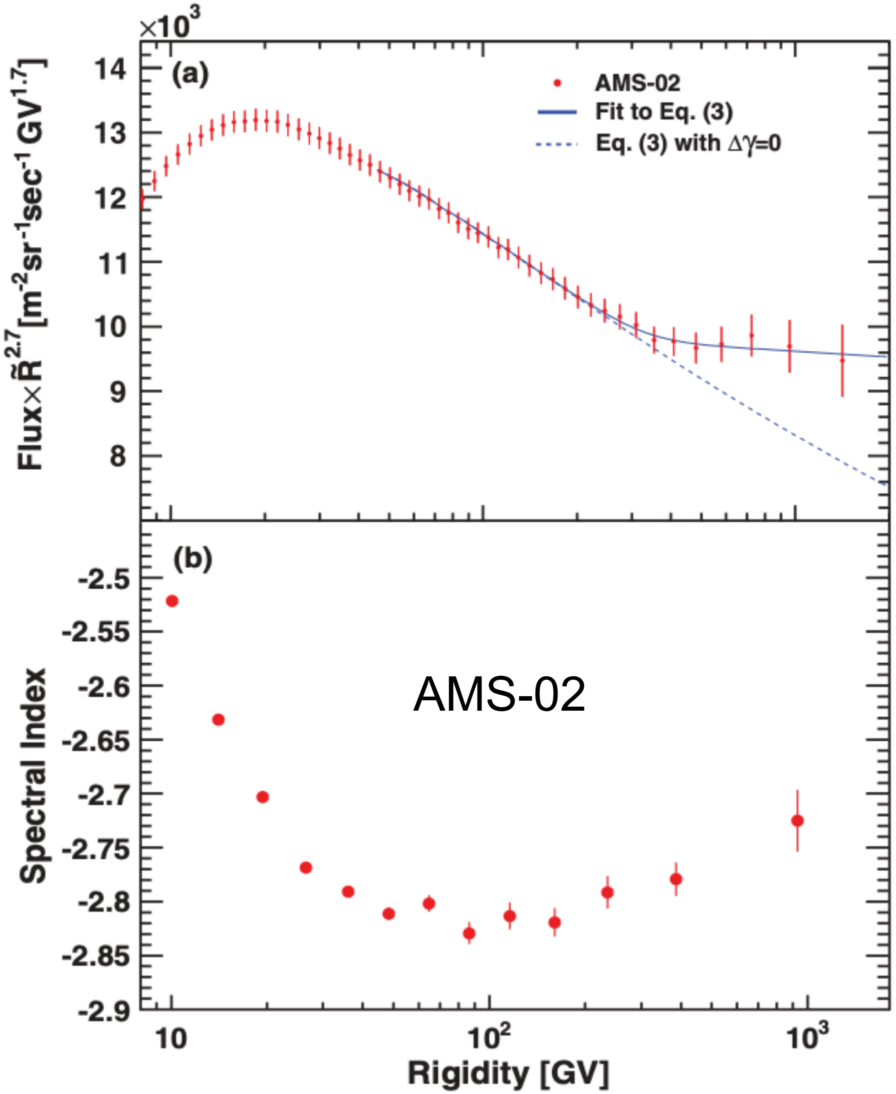
Consensus

- Single power law spectrum from 10 GeV (10^{10} eV) up to 1 PeV (10^{15} eV)
- Energy-dependent confinement in the Galactic halo
- Supernova remnants (SNR) as sources

Detection method

- Direct measurement (ballon, satellite or extensive air shower array), measure the local spectrum and anisotropy
- Indirect measurement (via Gamma-rays). spectrum and distribution in the Galaxy

Direct measurement



[Phys. Rev. Lett. 114, 171103 (2015)]

Direct measurement

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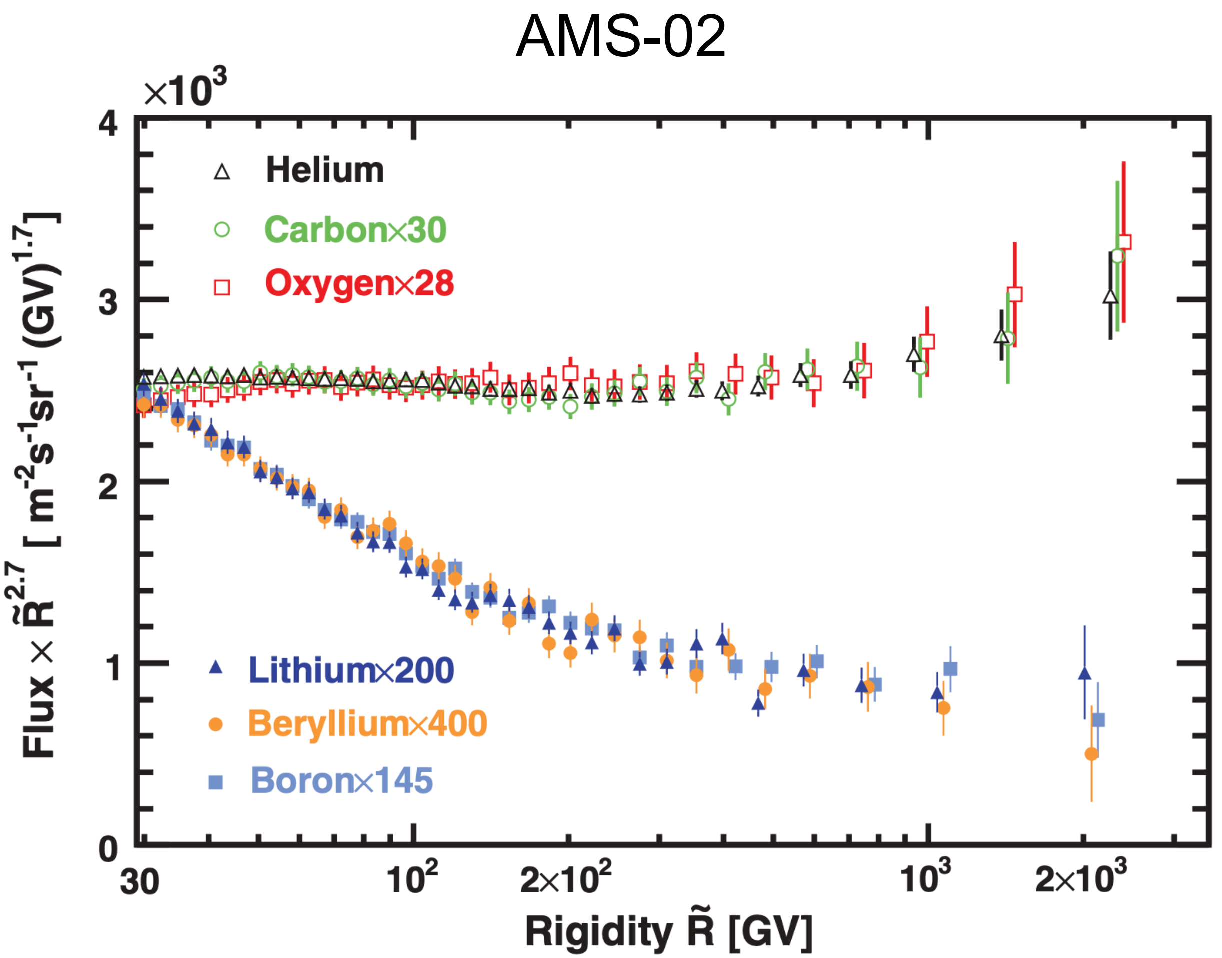

Phys. Rev. Lett. 120, 021101 (2018)

Direct measurement

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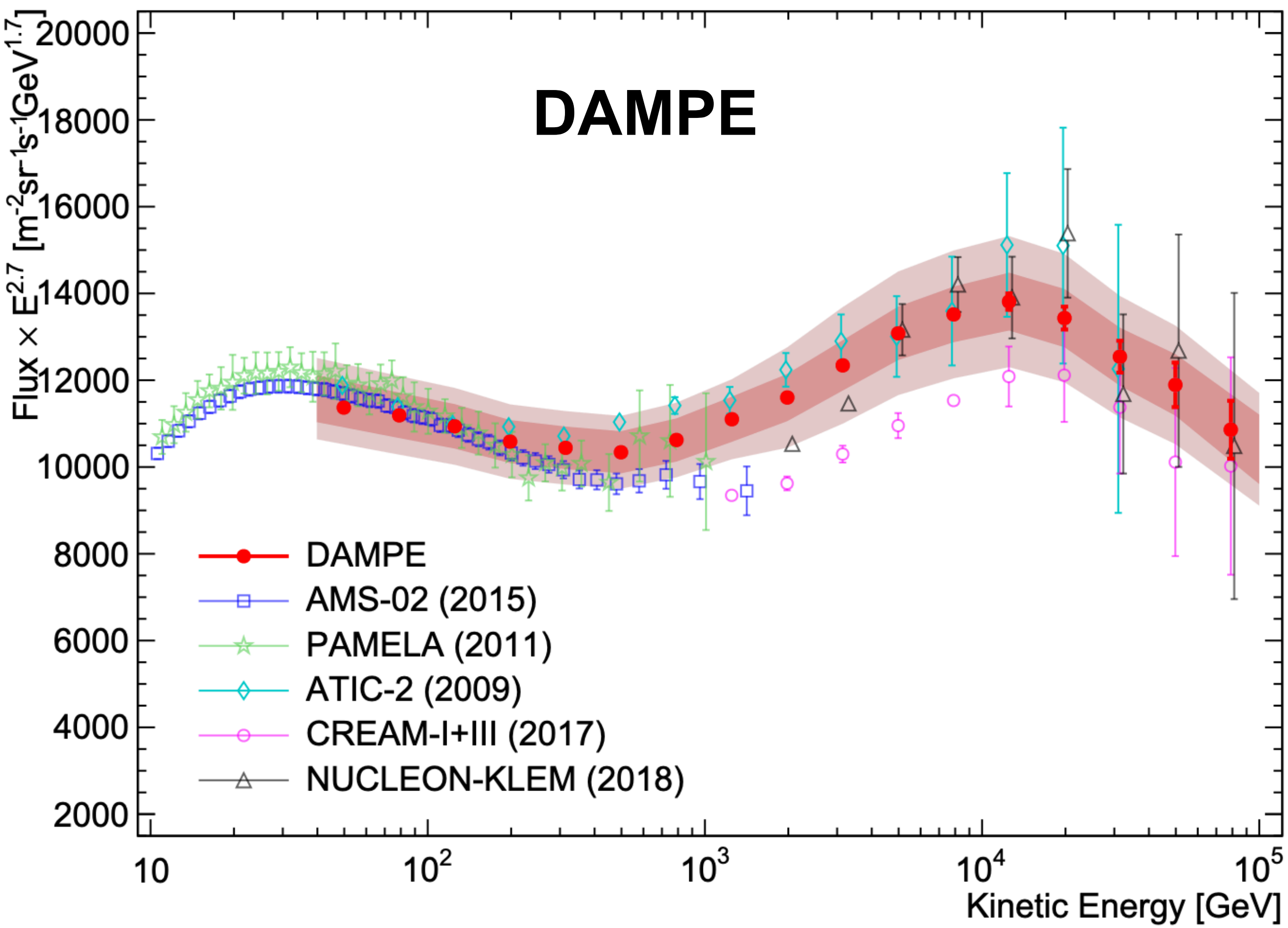
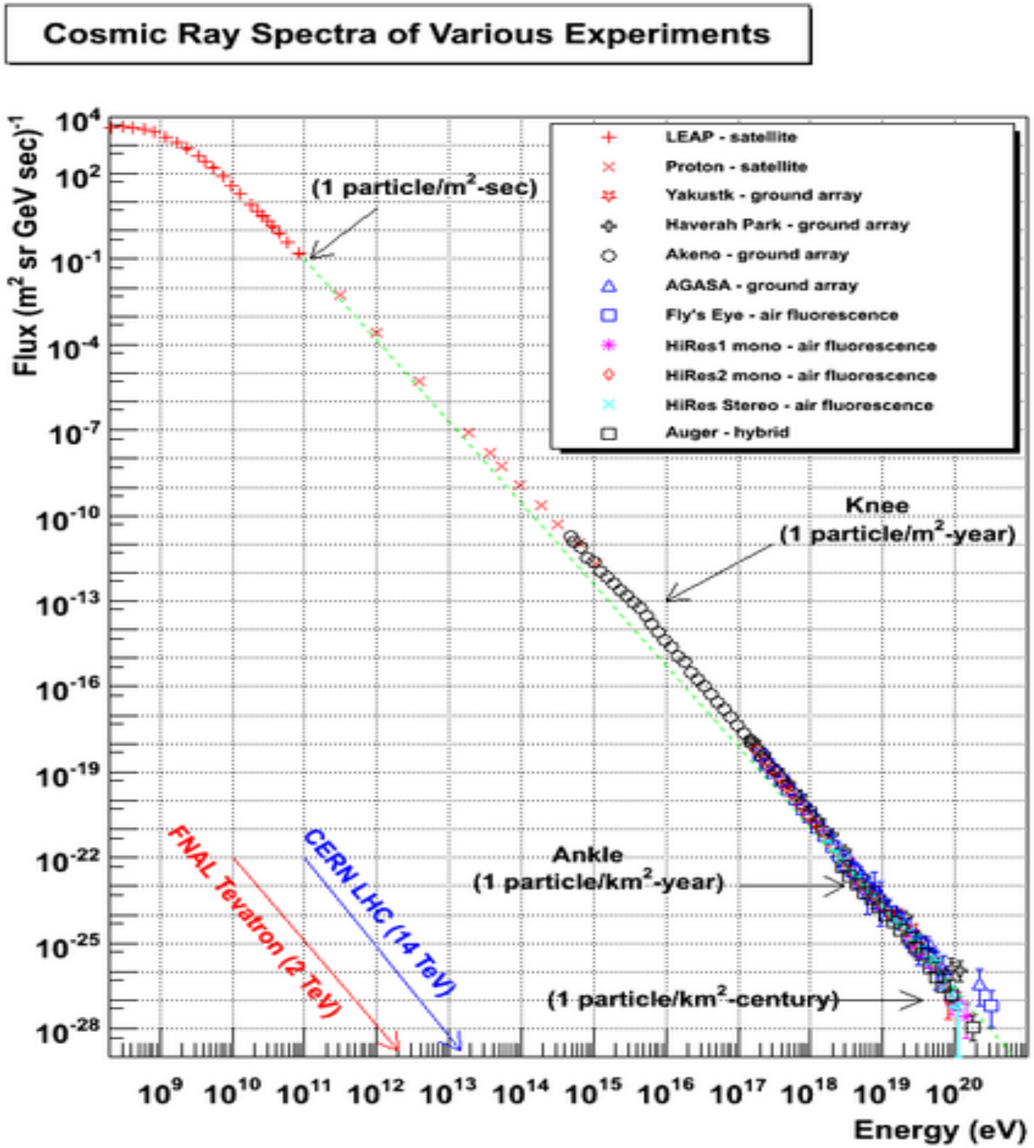


Science Advances, Vol. 5, no. 9 (2019)

Spectrum



https://www.physics.utah.edu/~whanlon/spectrum.html

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Spatial distribution from gamma-ray astronomy

Gamma-ray Astronomy

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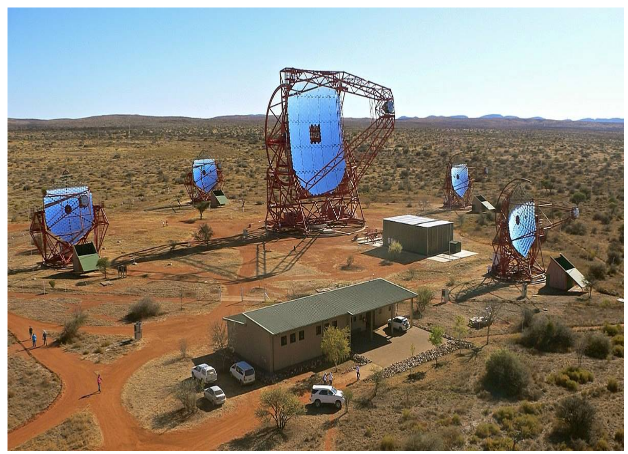
- Atmosphere is opaque for gamma-rays
- Satellite or air shower
- above 100 MeV

Fermi LAT
(0.1 ~ 1000 GeV)

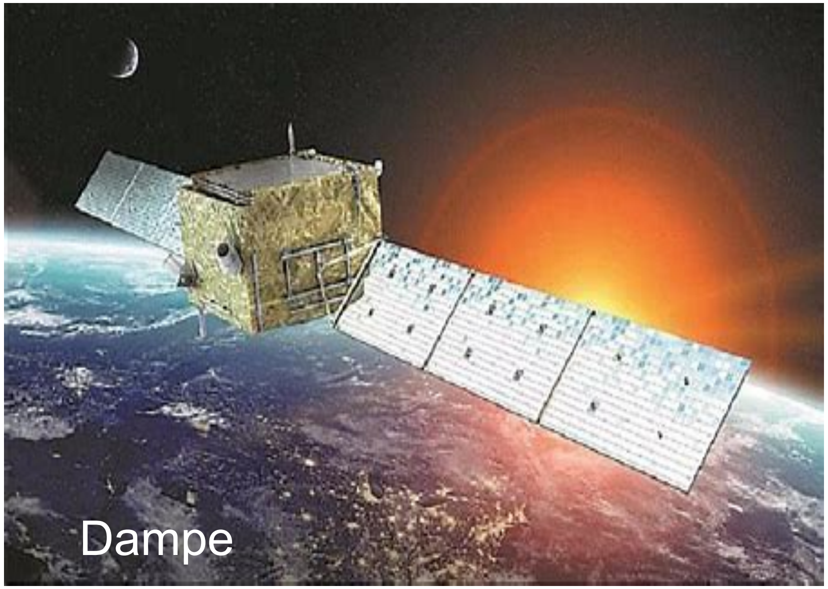


H.E.S.S
(0.1 ~ 10 TeV)

LHAASO
(>1 TeV)

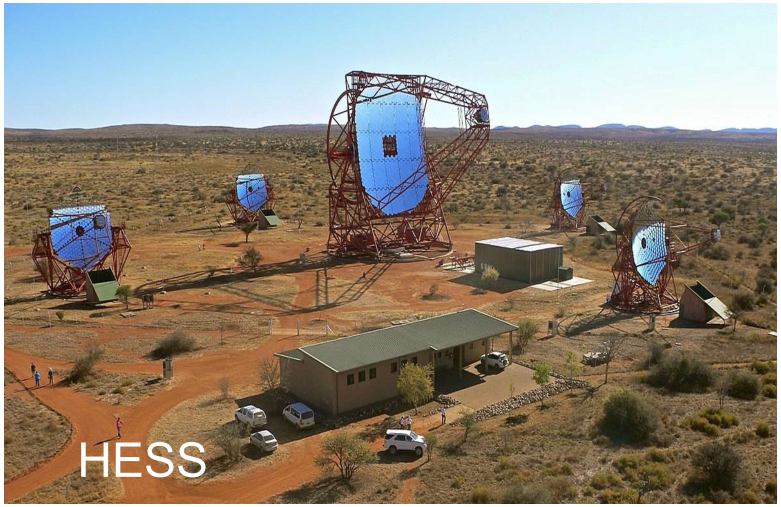
Spaceborne detectors

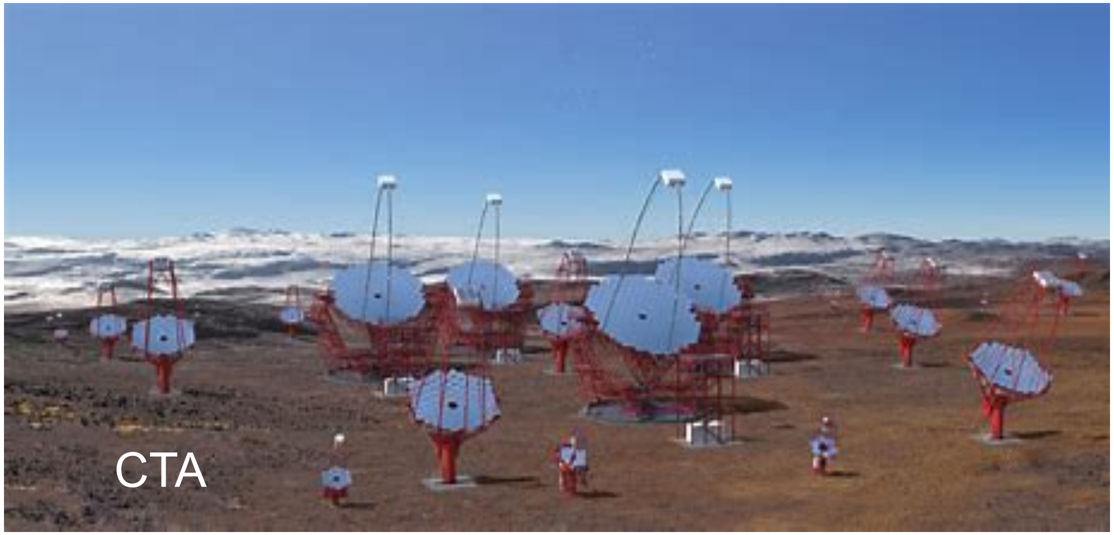
- pair conversion telescopes
- Large FOV (~ 3 sr)
- 0.1~ 1000 GeV
- small effective area (~ 1 m^2)
- moderate PSF (0.1 - 1 degree)
- continuous monitoring



Air Cherenkov telescope arrays

- detect cherenkov light of secondary electrons in air shower.
- small FOV (~ 5 degree)
- 0.1~ 100 TeV
- large effective area (~ 1e5 m^2)
- excellent PSF (down to 1 arcmin)
- only at clear night without moon





Extensive air showers

- detect muon/electrons in air shower (scintillator/water cherenkov).
- large FOV (~ several sr)
- above 1 TeV
- large effective area (~ 1e5 m^2)
- poor PSF (0.3 degree)
- continuous monitoring

From gamma-rays

- Gamma-ray emission (in molecular clouds or diffuse):

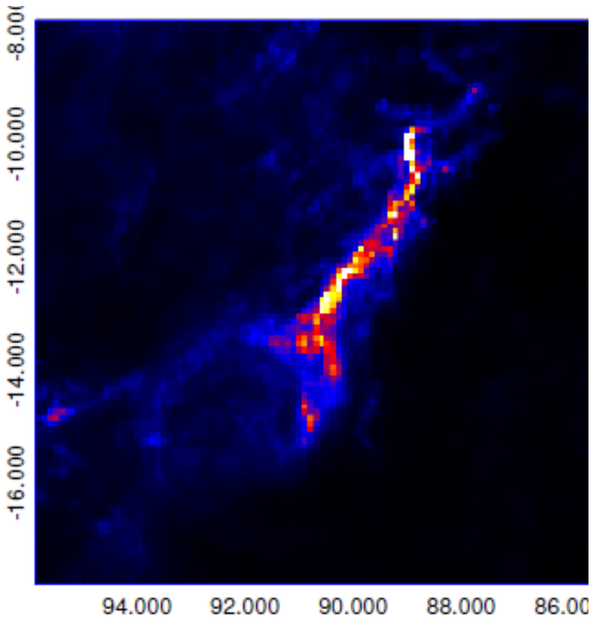
Point sources+ CR interaction with ambient gas + ICs +isotropic

- CR interaction with gas dominates in dense environment.

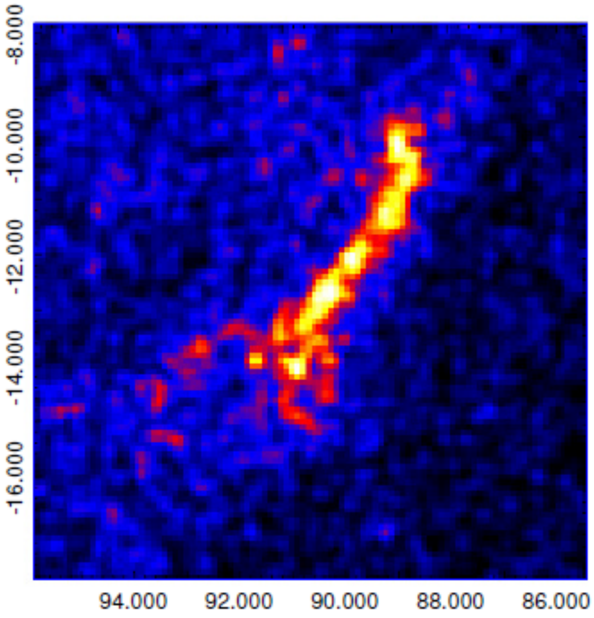
- Gamma-ray map + gas distribution -> CR distribution

gamma-rays from giant molecular clouds (GMCs)

- Gamma-rays show good correlation with gas (CR uniformly distributed inside GMCs)
- Can be used to study the CR spectra



Gas (CO) distribution



gamma-ray observations (GeV)

Derived CR spectrum

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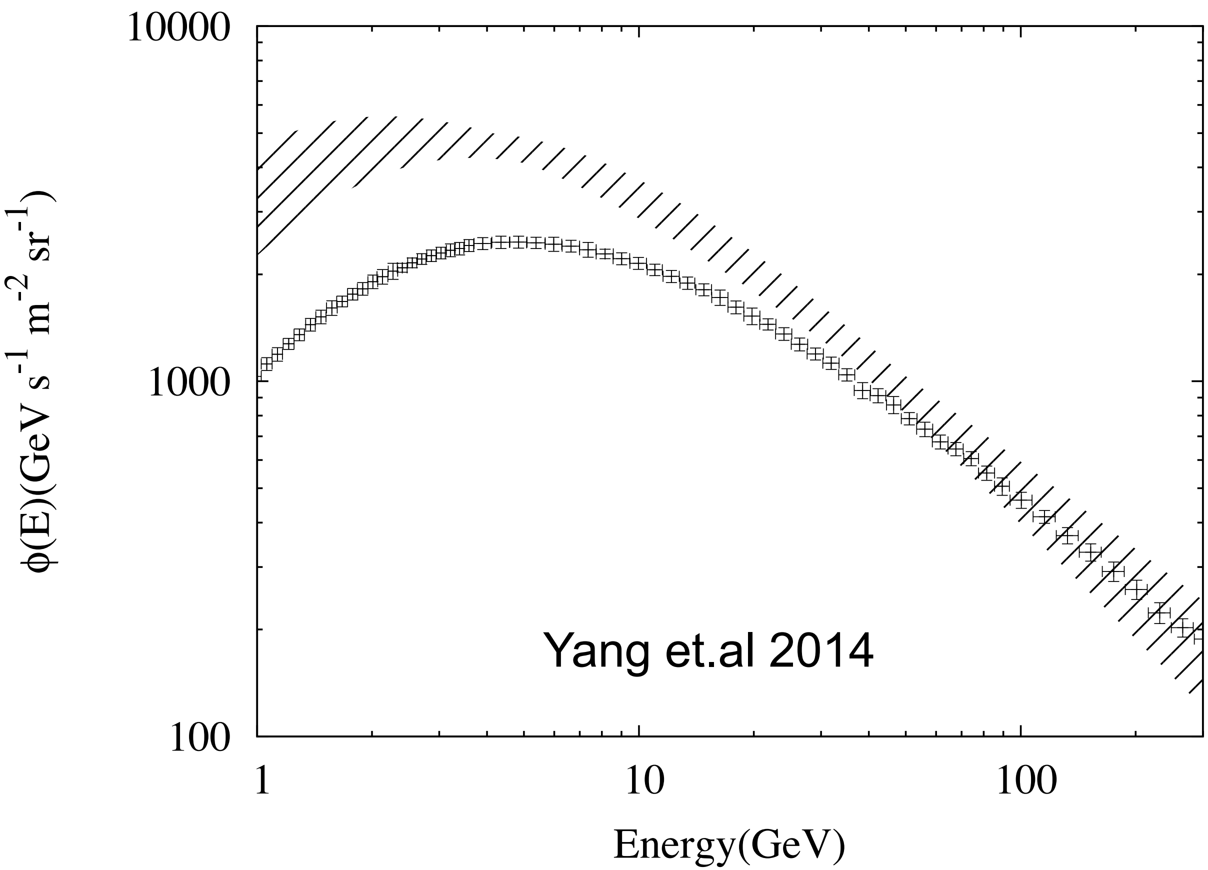

In comparison with the Local Measured CR: consistent above 10 GeV (solar modulation)

uniform or not?

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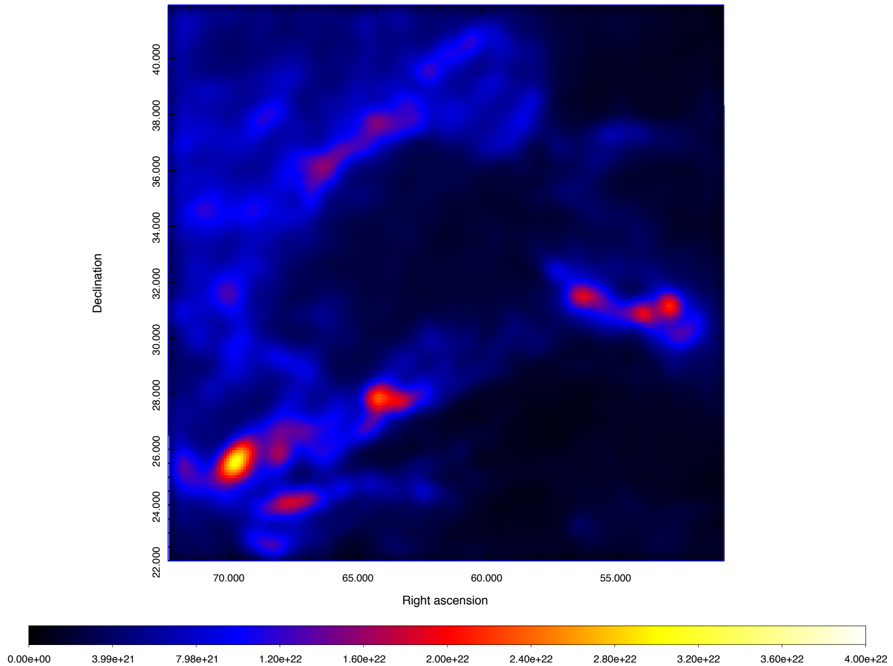
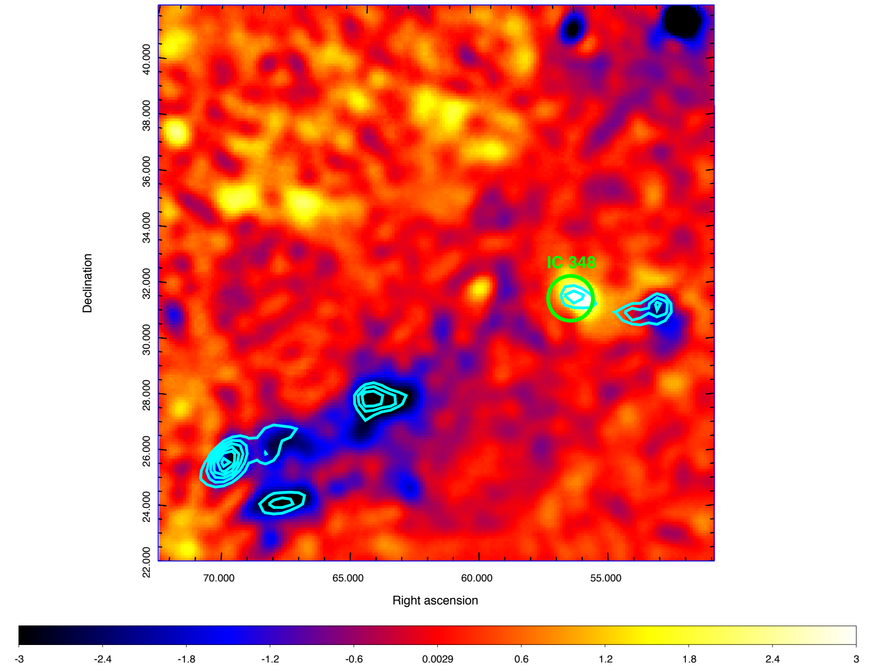
Test the uniformity of CRs

- Some hint of inhomogeneous distribution in Taurus-Perseus region



Dust opacity (gas distribution)



gamma residual (CR density)

Uniform or not?

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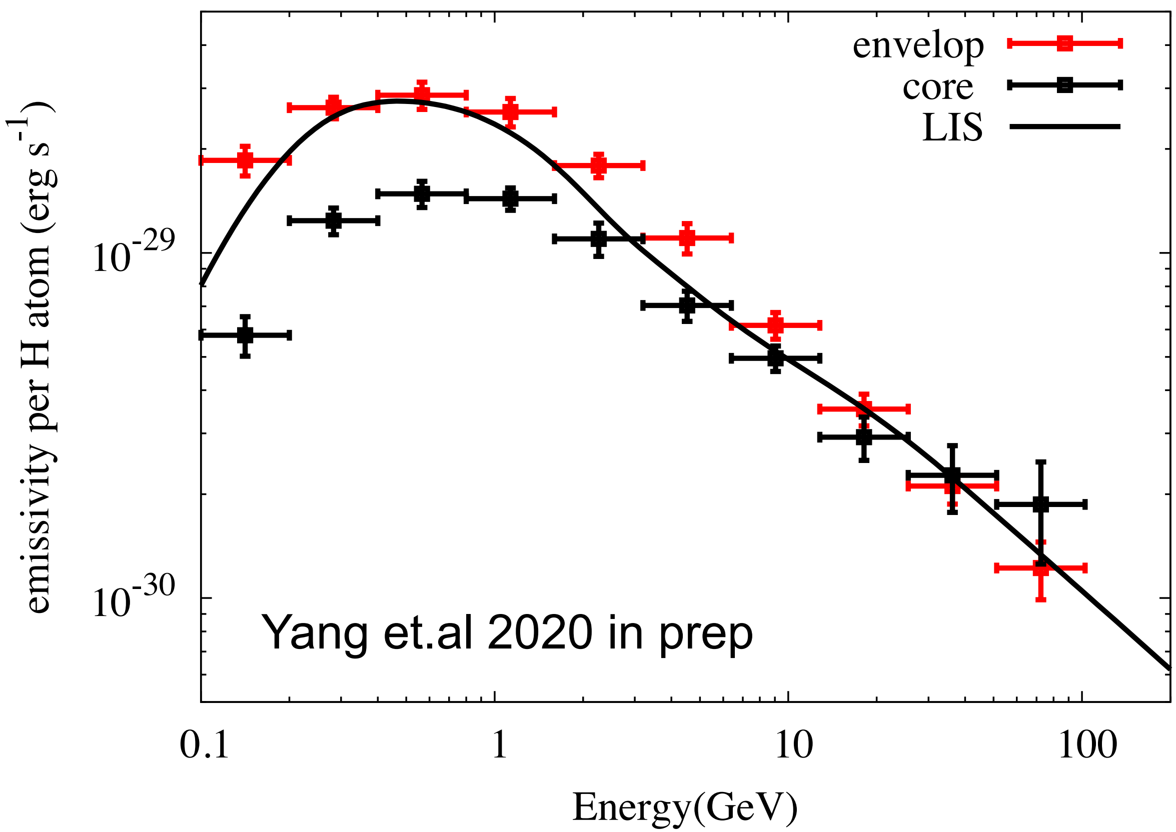
Test the uniformity of CRs

- Low energy CRs cannot penetrate into the core:
  slower diffusion due to higher turbulence inside GMCs?



Diffuse gamma-ray emission

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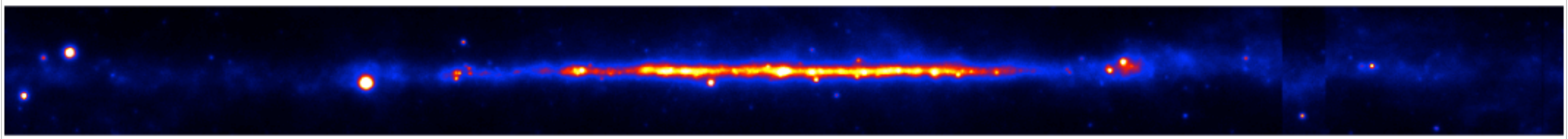
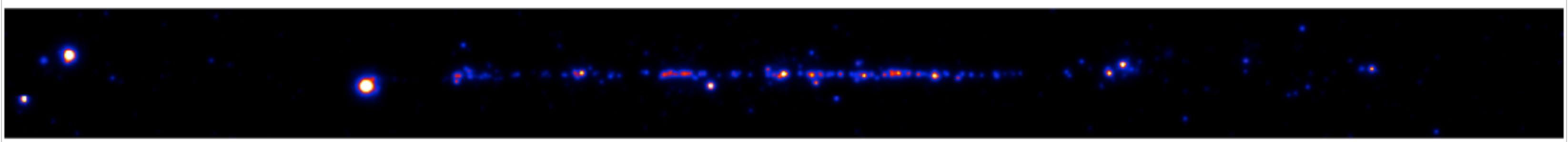
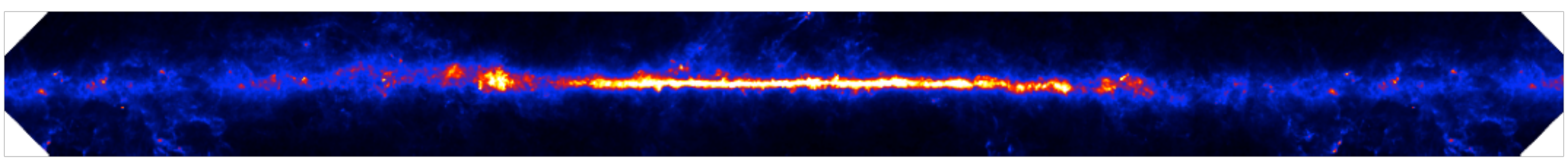
Gamma-ray counts map



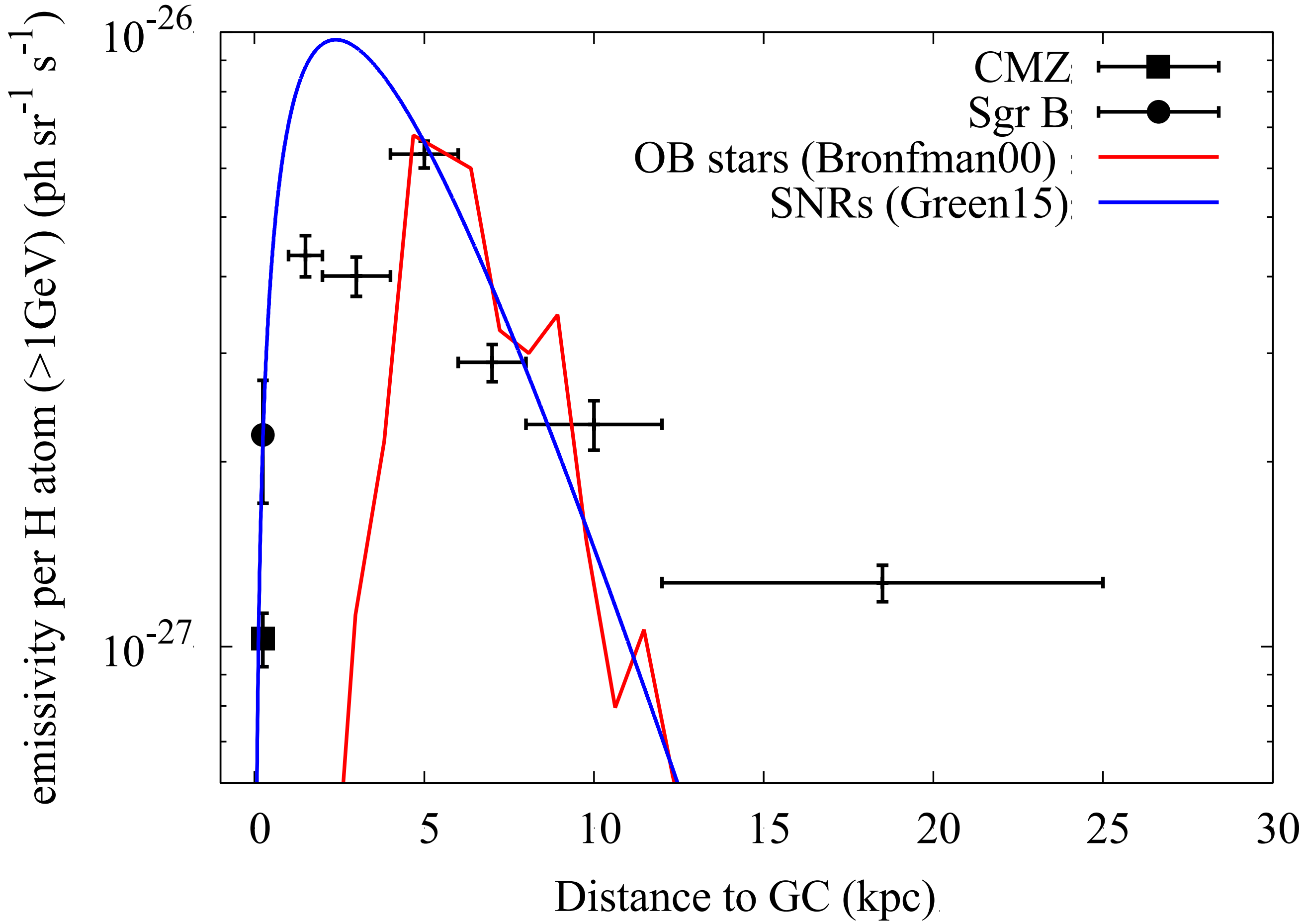
Point source contribution



Dust opacity map (gas column)



CR Radial distributions



Yang et.al 2016

Hardening towards GC

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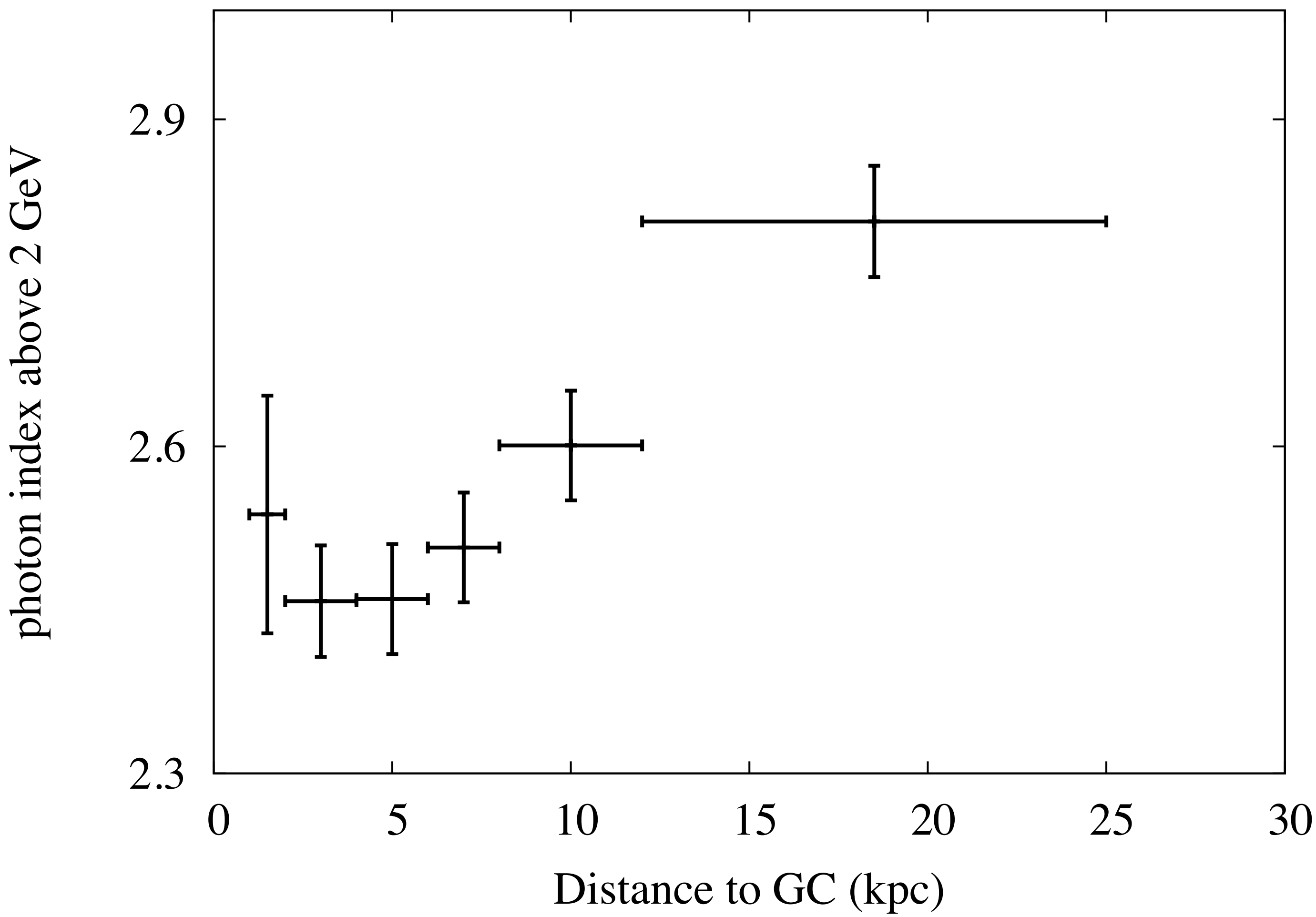

Yang et.al 2016

More GMCs!

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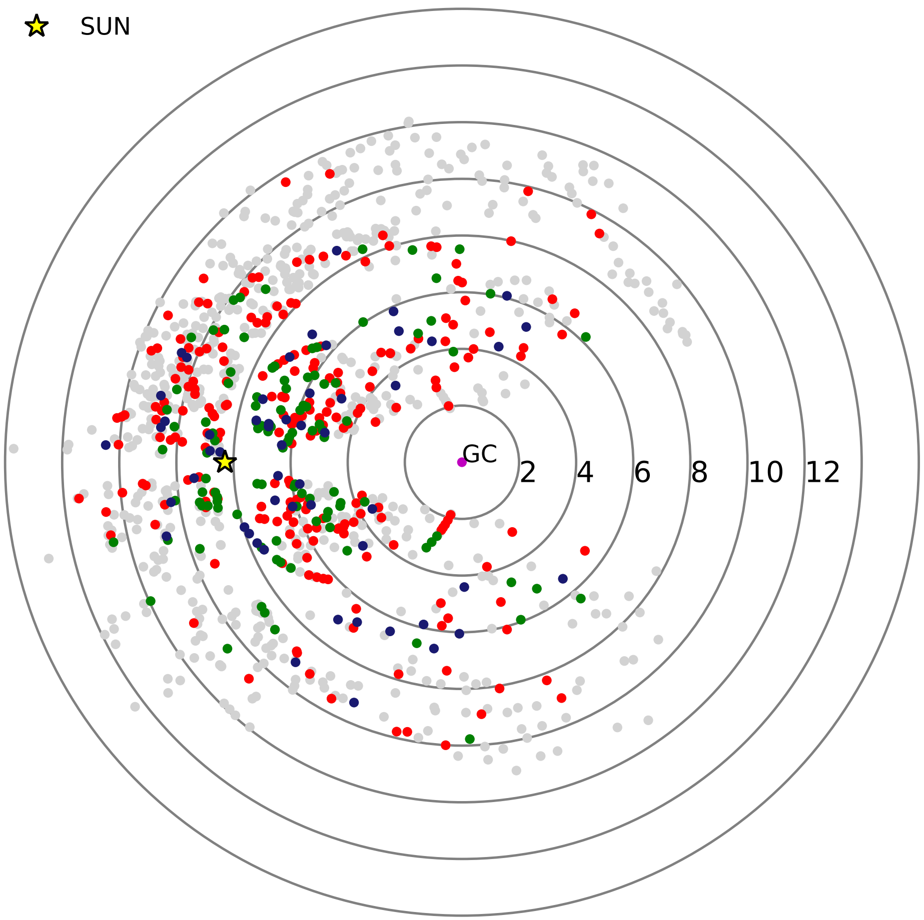

Aharonian et.al 2019

- Rice et.al (2016) have identified thousands of Molecular Clouds in the Galaxy
- Possible to measure CR density in each position of the Galaxy.

More GMCs

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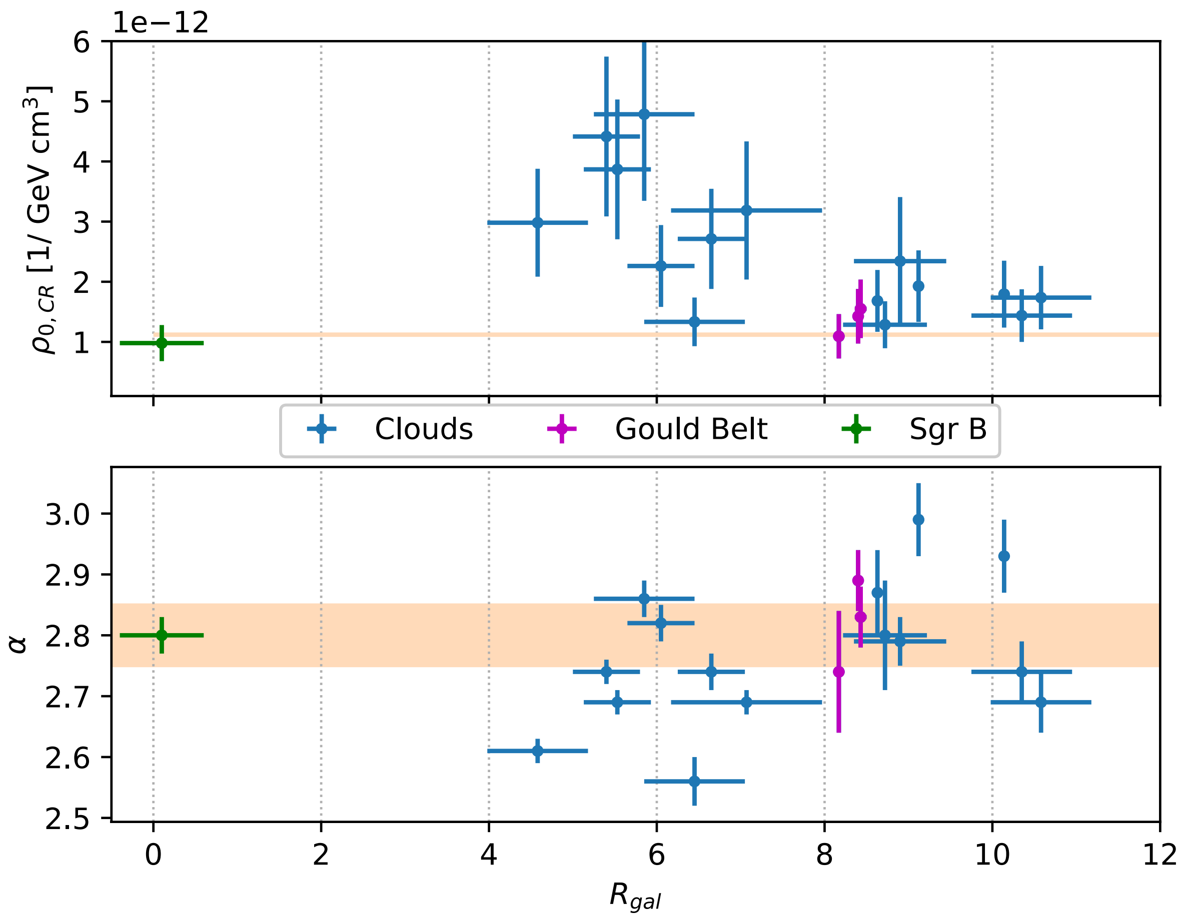

Aharonian et.al 2019

- The enhancement and hardening is caused due to the CR sources?
- A uniform CR “sea” plus some “islands” with higher density and harder spectra?

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LECRs

(Yang, Liu and Aharonian 2020 in prep)

Low energy (LE) CRs

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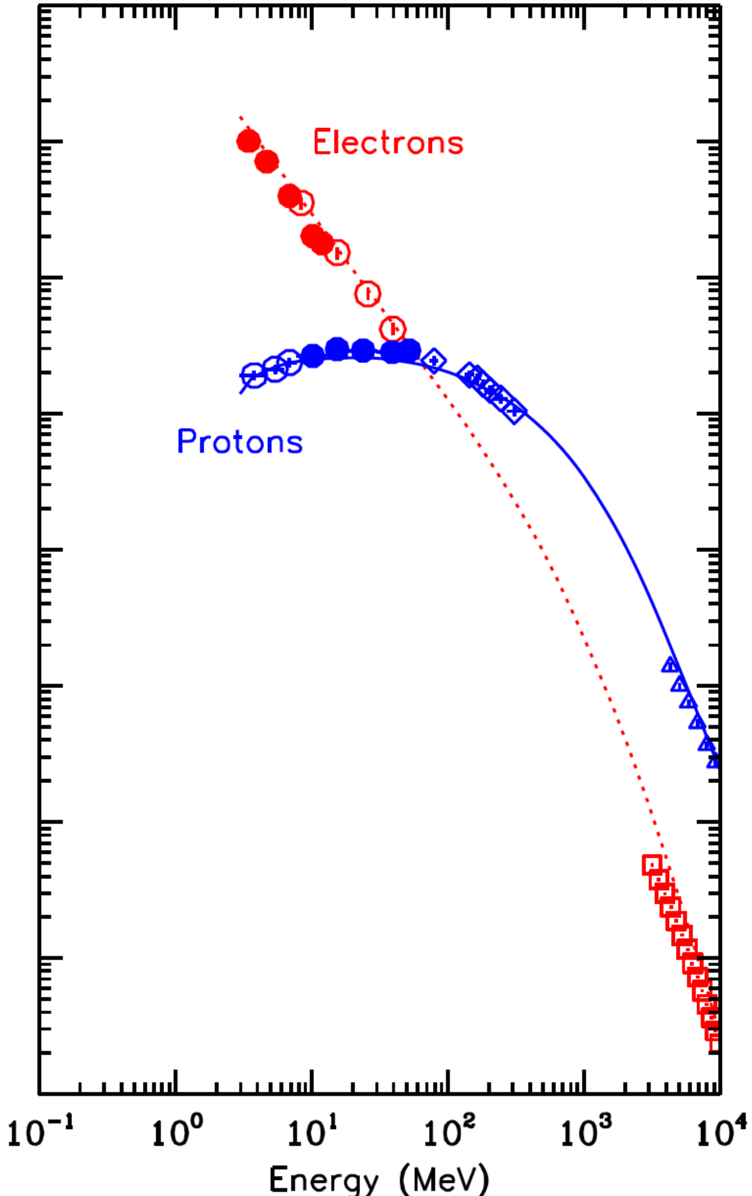
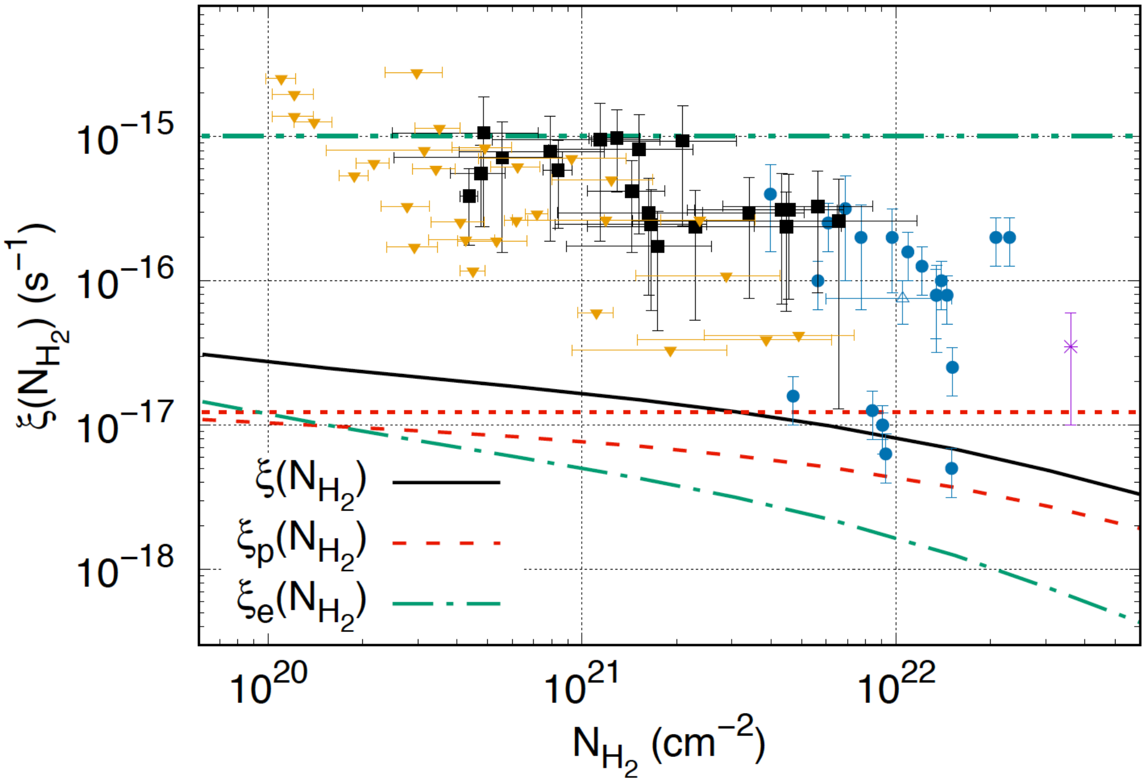
- E < 100 MeV, No pion-decay gamma-rays
- significant contribution to the energy density of ISM (~ eV/cm^3)
- Heating the gas
- Govern the astro-chemistry
- dominate ionization in MCs
- At MeV energy ionization dominate cooling
- Voyager measurement in ISM (Cummings et.al 2016)



LECRs: Ionization

- CR dominates ionization inside MC cores (UV shielded)
- The measured Ionization rates from astro-chemistry are larger than expected



Calculation from Phan et.al 2018, Black curve is the ionization rate assuming voyager measurement is the universal LECR spectrum

LE CR propagation

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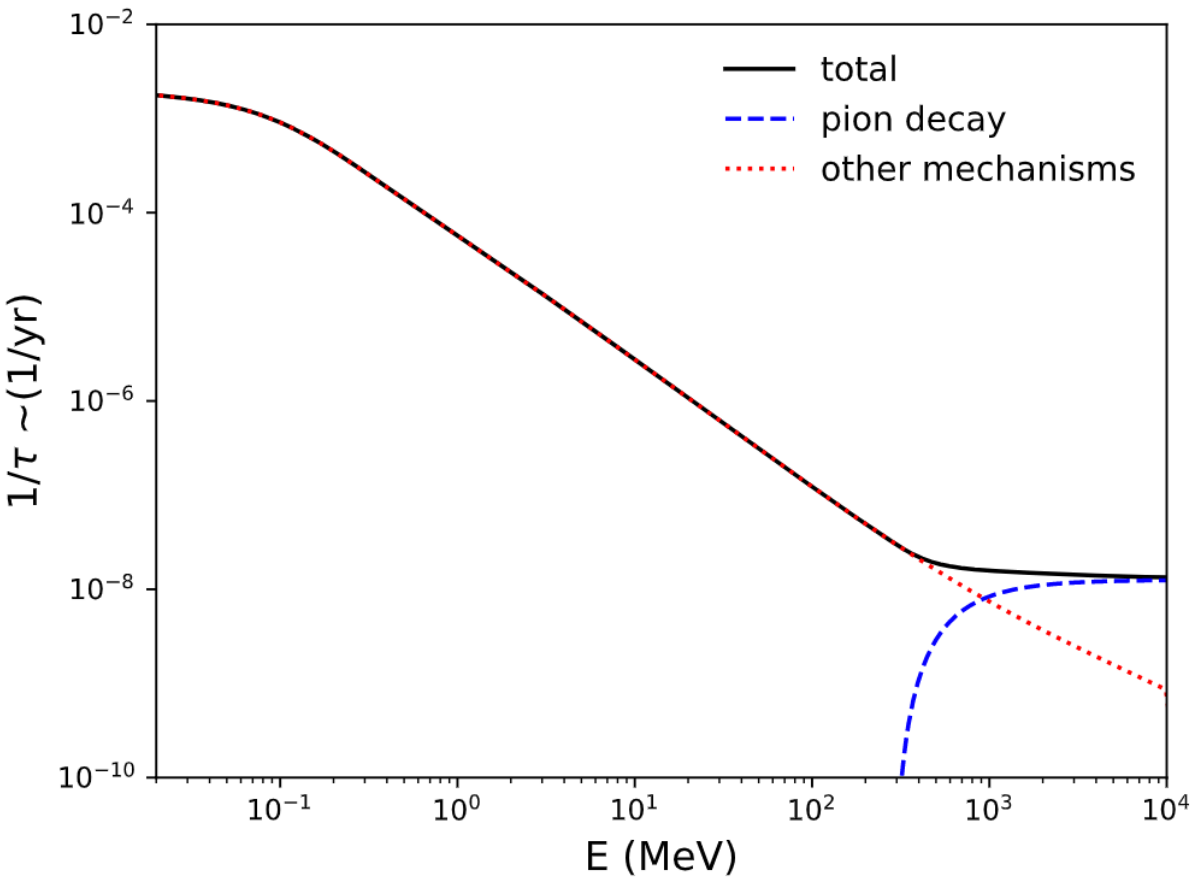
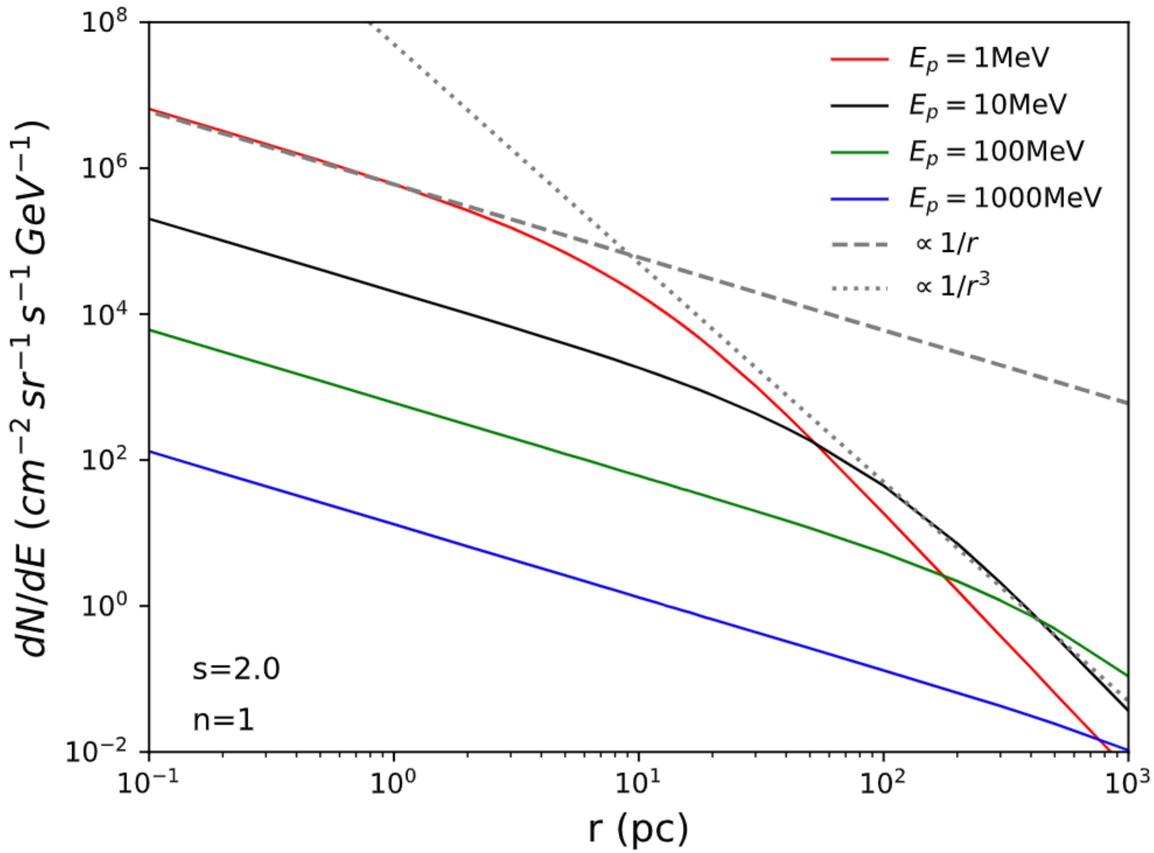
- But is the LECR spectrum universal?
- For LECR ionization cooling (see below) is significant in MeV range and the propagation is slow



LECR propagation

- LECR should be similar to VHE electrons, cannot propagate far
  Flux can be very different at different distances to the source



Gamma-ray line

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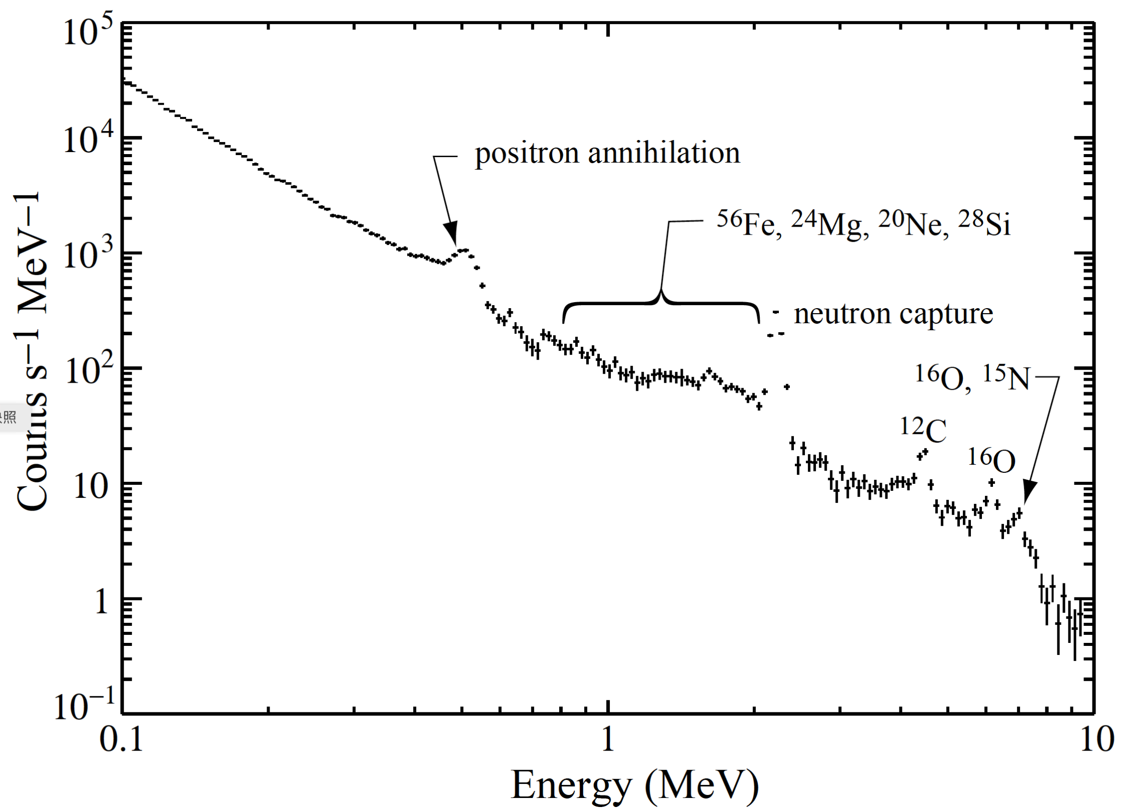
- The same CRs can be studied also in gamma-rays through de-excitation line of nuclei
- Well studied in solar flare (Kozlovsky et.al 2002)



Gamma-ray line

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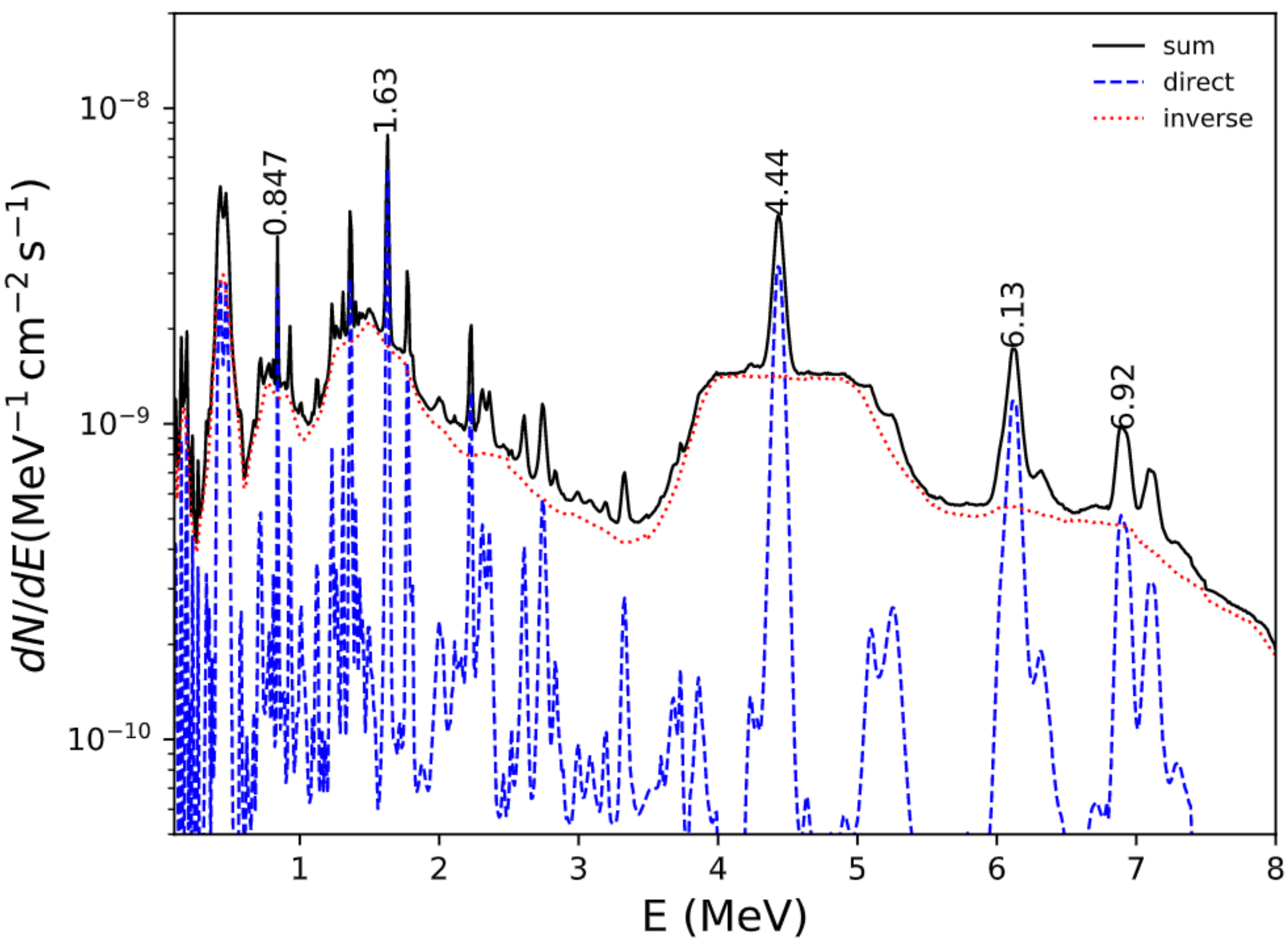
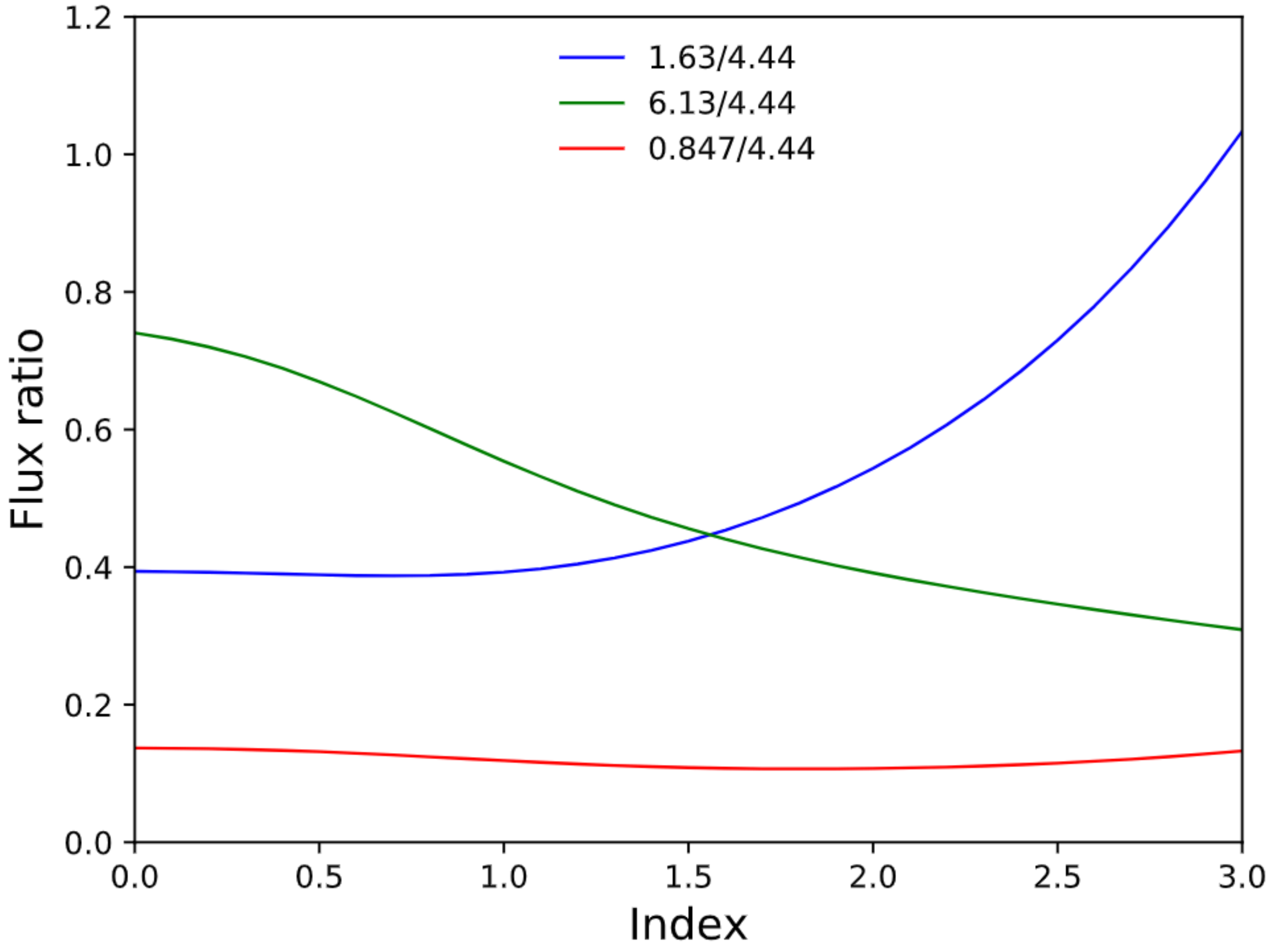


Inverse and direct process

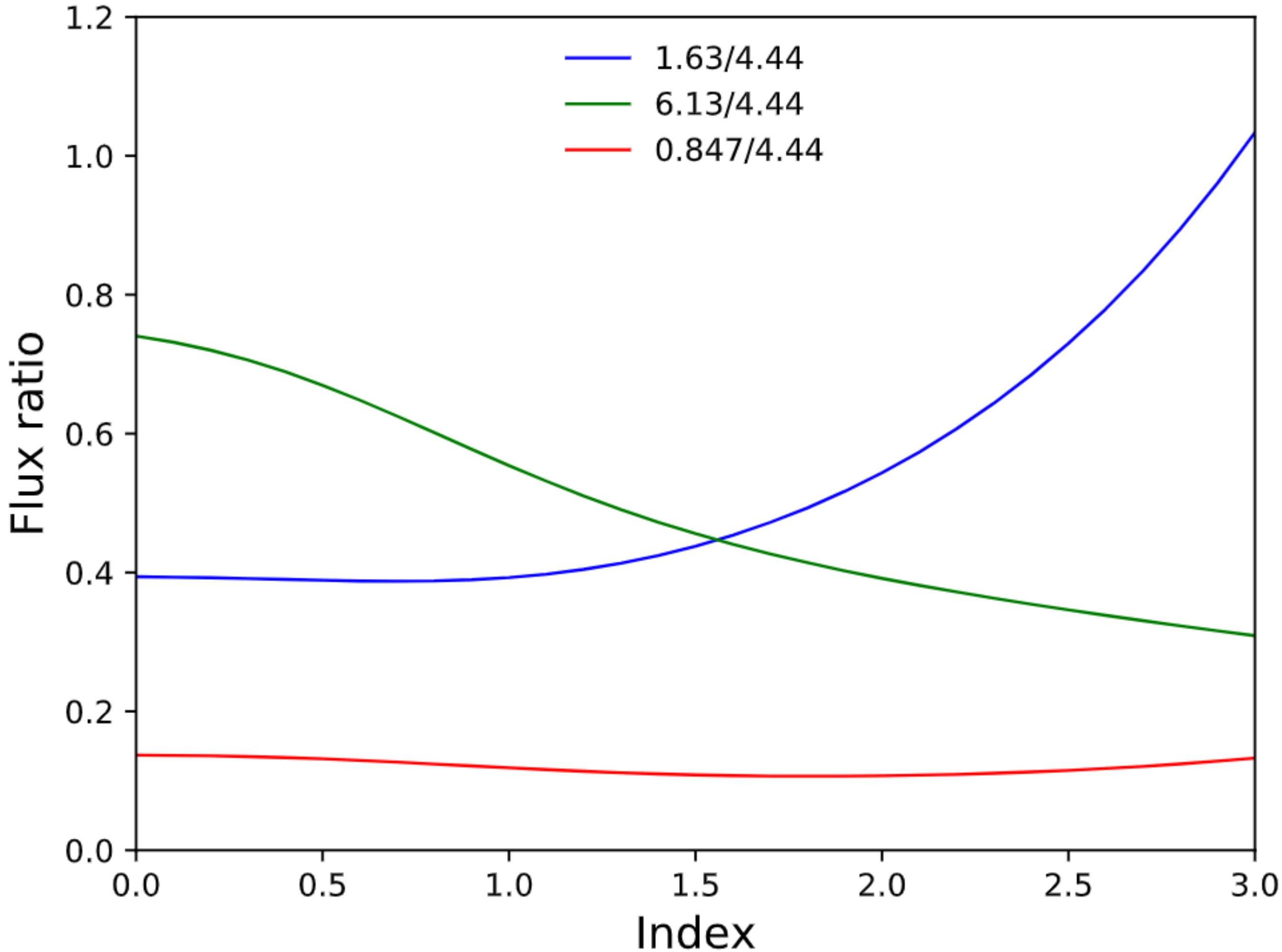
line ratios

Use line ratio to diagnose CR spectrum



line ratios

Use line ratio to diagnose CR spectrum



MeV gamma-ray, “LAST” electromagnetic window and interesting physics

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PeVatrons

Spectrum

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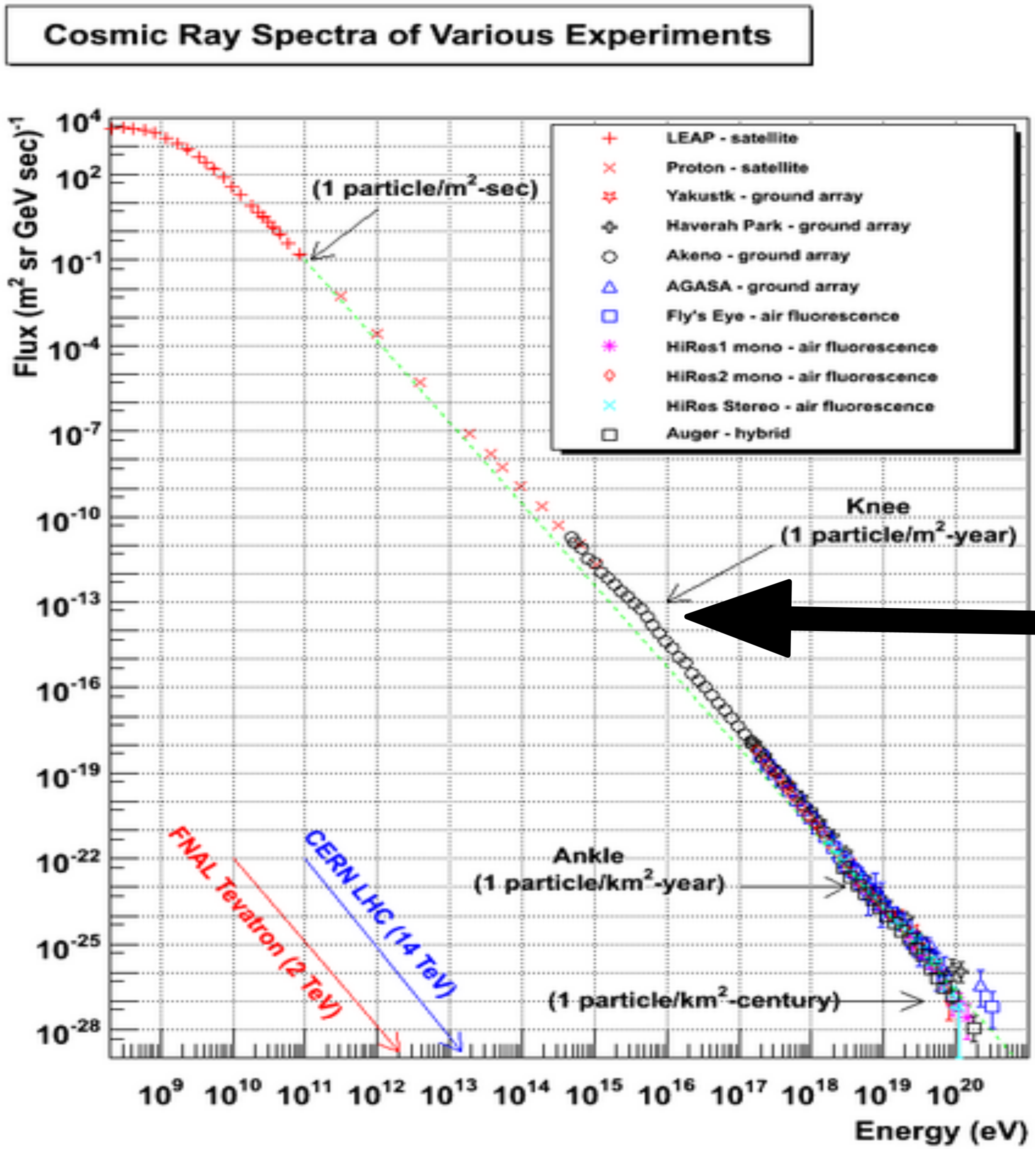


Knee: GCR at least to PeV

https://www.physics.utah.edu/~whanlon/spectrum.html

Galactic Cosmic Rays (GCR)

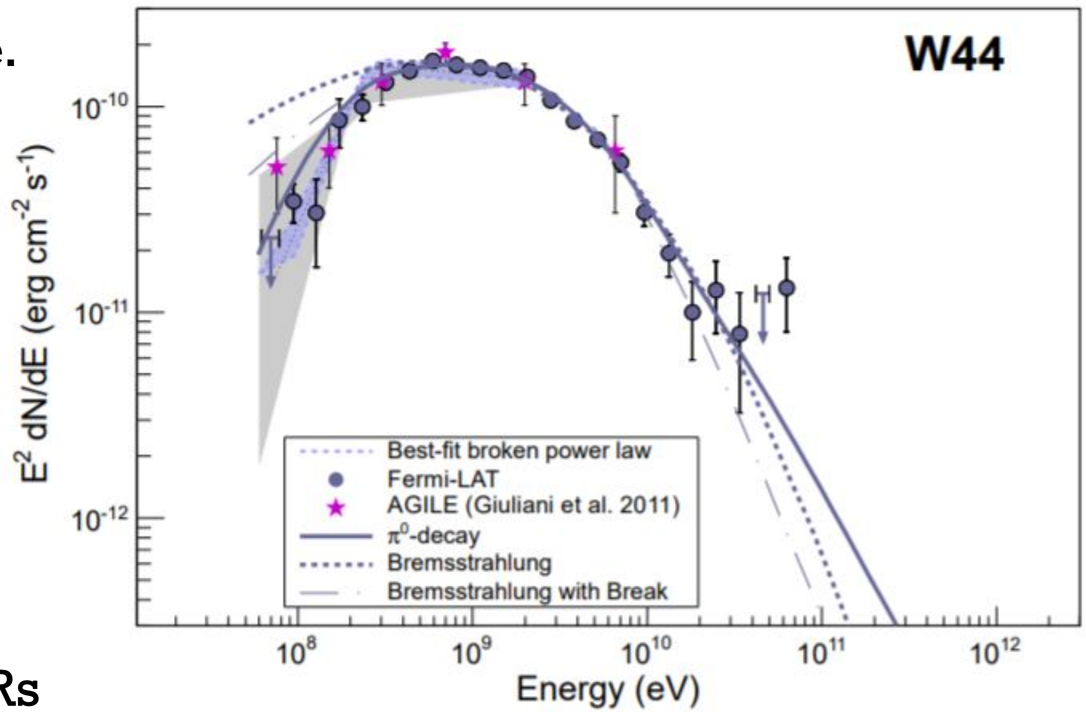
Why SNRs?

- Energy budget reasonable: 1e40 erg/s considering 10% efficiency
- Acceleration: 1st order Fermi acceleration in the shock front
- Observation proofs

SNRs as CR source?

Mid-age SNRs

- Clear Pion-decay feature.
- Hadronic origin or Bremsstrahlung ?
- Break at ~ 10 GeV
- Cannot account for all CRs up to PeV



Fermi Collaboration 2013

Gamma-ray observation of Young SNRs

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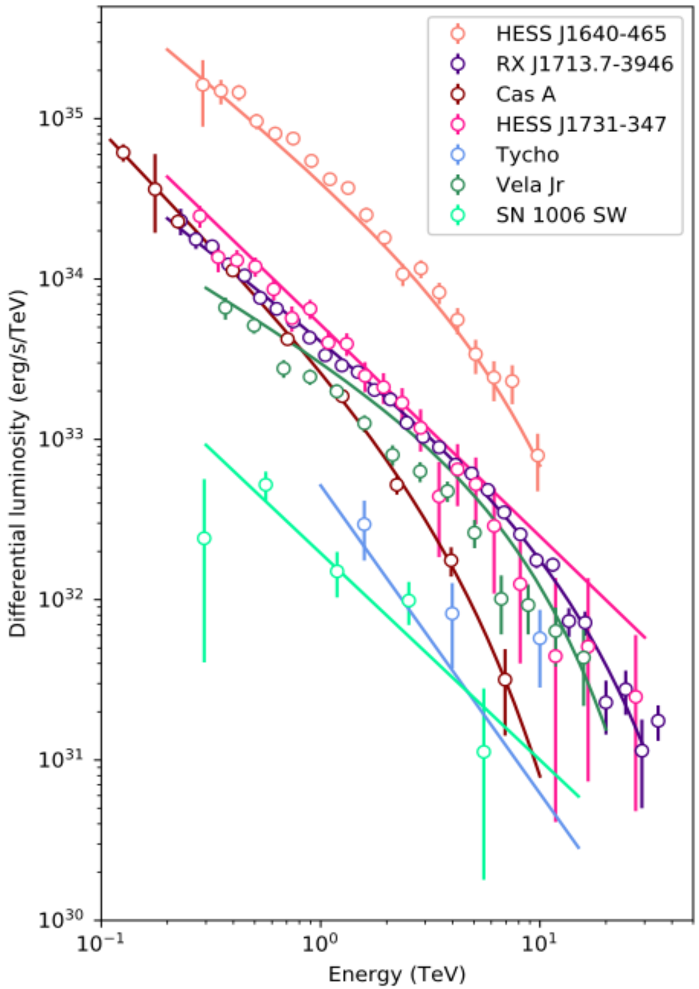

- All gamma-ray spectrum of young SNRs shows soft spectrum or early cutoff at ~ 10 TeV
- corresponding to CR energy of 100 TeV
- Hard to address a single power law spectrum of CRs up to PeV

Very young SNRs?

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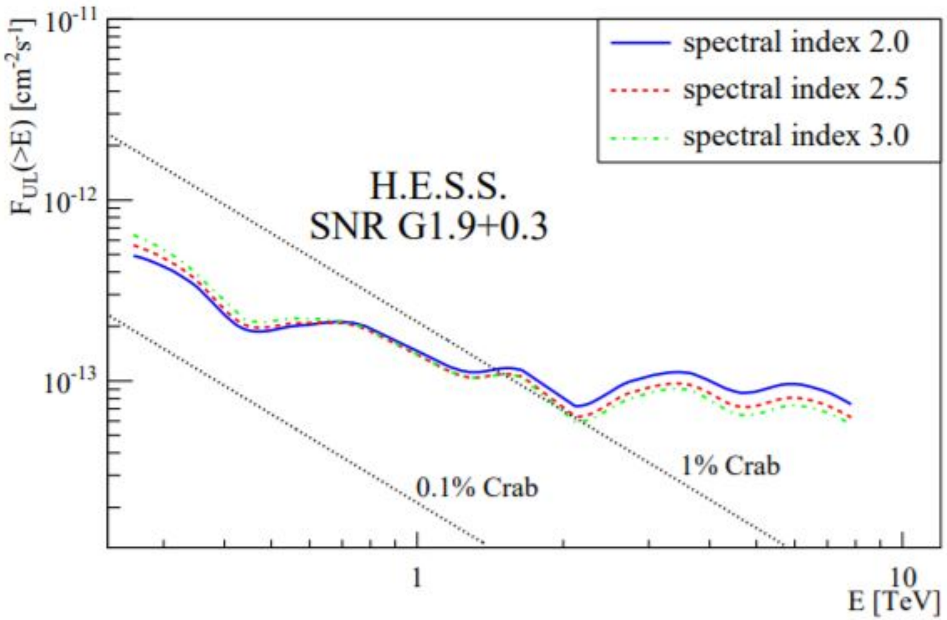
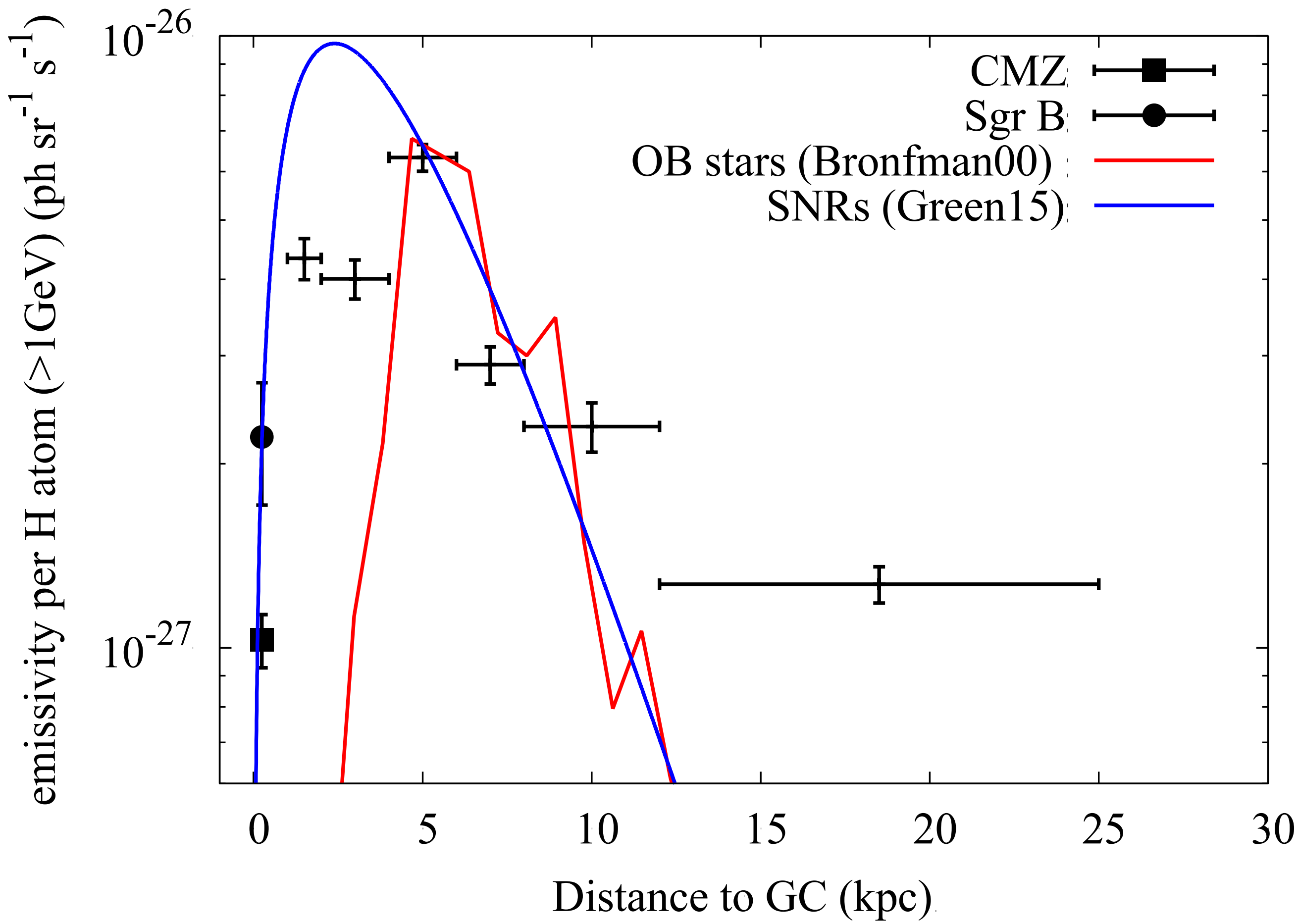
- PeVatron phase could be accomplished only during the first years of the explosion (e.g., Bell et.al 2013)
- The youngest SNR in the Galaxy: G1.9+0.3, t ~ 100 yr
- VHE protons cannot propagate more than 30 pc.
- HESS reveals L(>1 TeV) < 1e32 erg/s can be used to set limit on proton energy budget.
- Considering a high density in the vicinity (near GC), the total energy on VHE protons are below 1e45 erg. Not enough to account for the CR flux up to the knee.



Alternative sources： Young massive clusters

- Isotope measurement favor a superbubble origin. (W.R Binns 2016)
- Most of OB stars exist in associations or clusters, stellar wind can accelerate CRs (Cesarsky & Montmerle 83).
- Efficiency may even better than SNR (high speed wind lasts much longer than SNR shock)
- Sufficient wind power (10^{38} - 10^{39} erg/s for each cluster, more than -10^{41} erg/s in the Galaxy) to account for CRs

CR Radial distributions



Yang et.al 2016

Young massive clusters

Westerlund 2 (HST image)

NGC 3603 (VLT image)

- More than dozens of OB stars and WRs
- Compact structures (~ pc)

Young massive clusters

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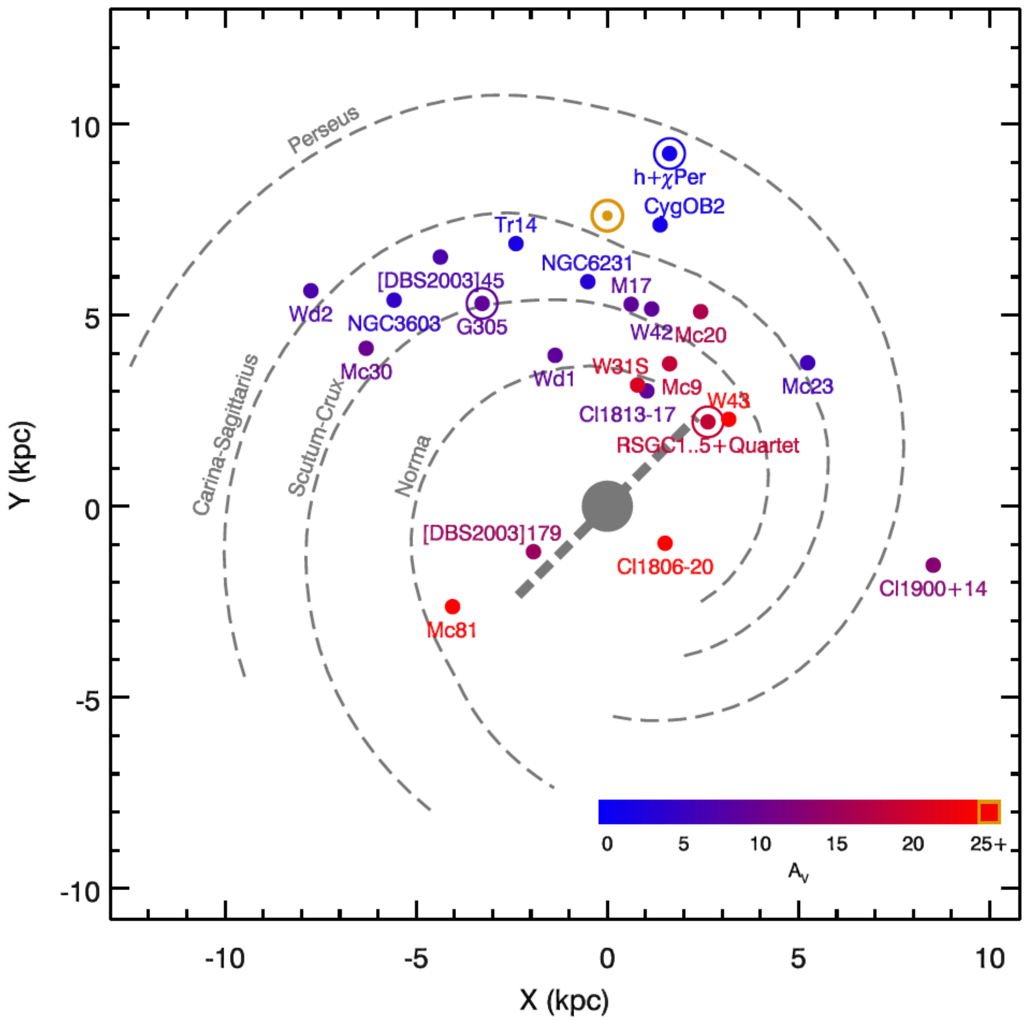

Davies et.al 2011

- ~20 in our Galaxy
- More to be discovered (high extinction in Galactic plane)

Stellar type	$\log[\dot{M}]$ $M_\odot\ yr^{-1}$	V_∞ [km s^{-1}]
WNL	-4.2	1650
WNE	-4.5	1900
WC6-9	-4.4	1800
WC4-5	-4.7	2800
WO	-5.0	3500
O3	-5.2	3190
O4	-5.4	2950
O4.5	-5.5	2900
O5	-5.6	2875

- The wind power of a single young star can be as high as 1e37 erg/s

Acceleration site?

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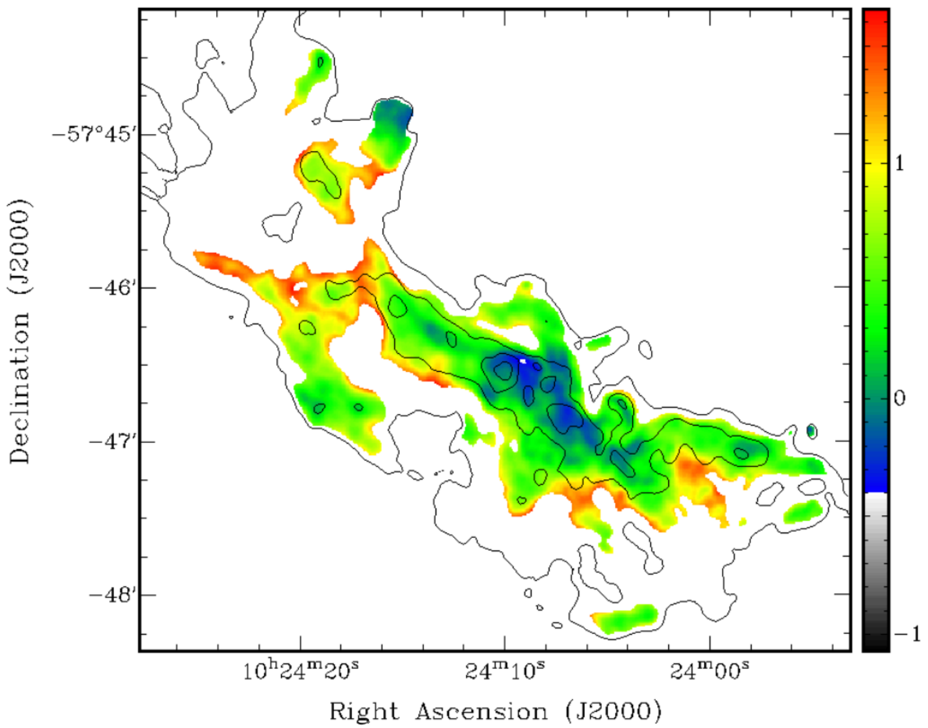
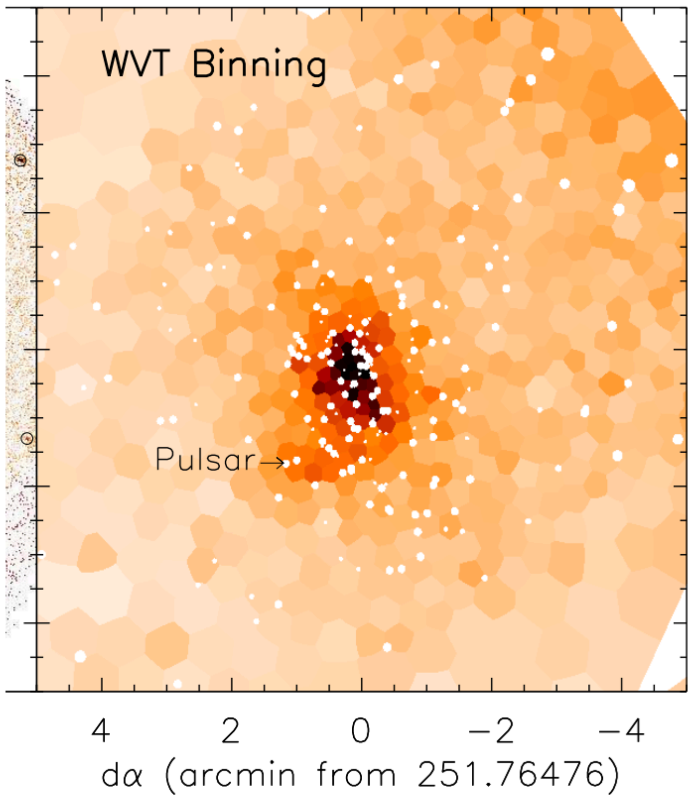


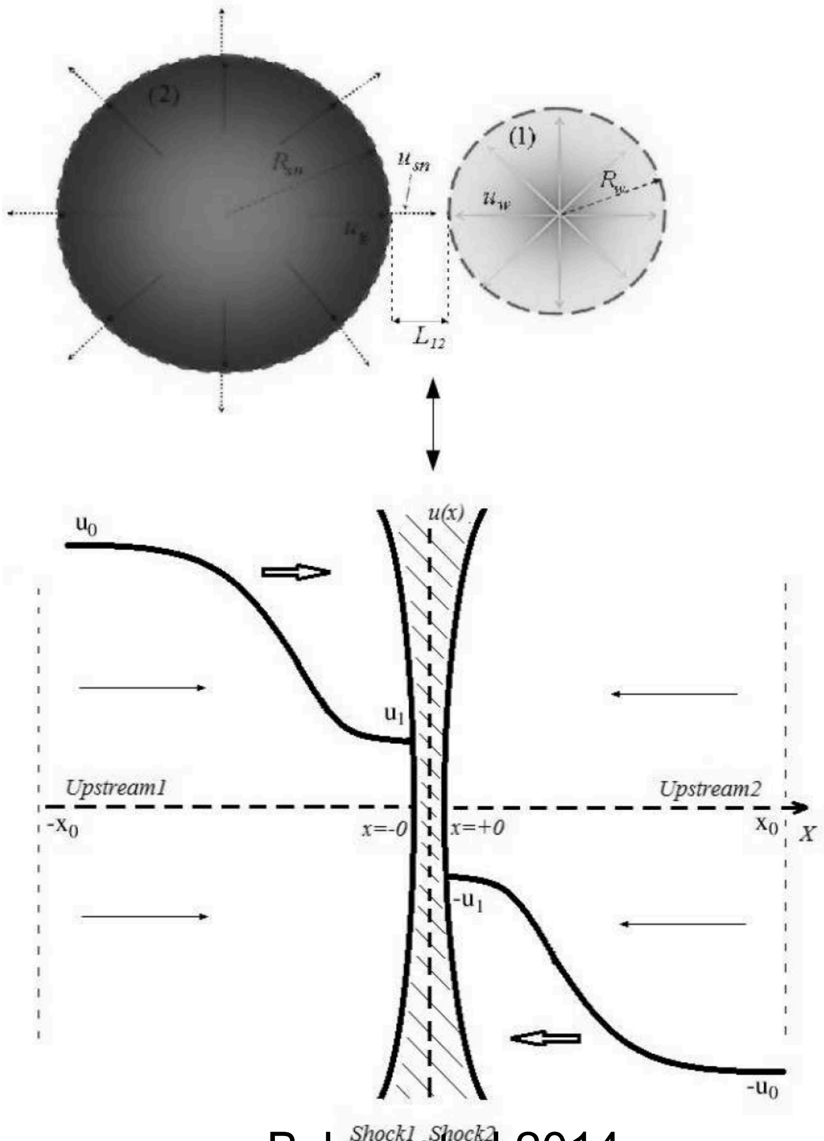
Radio continuum of Westerlund 2
(ACTA 9GHz)



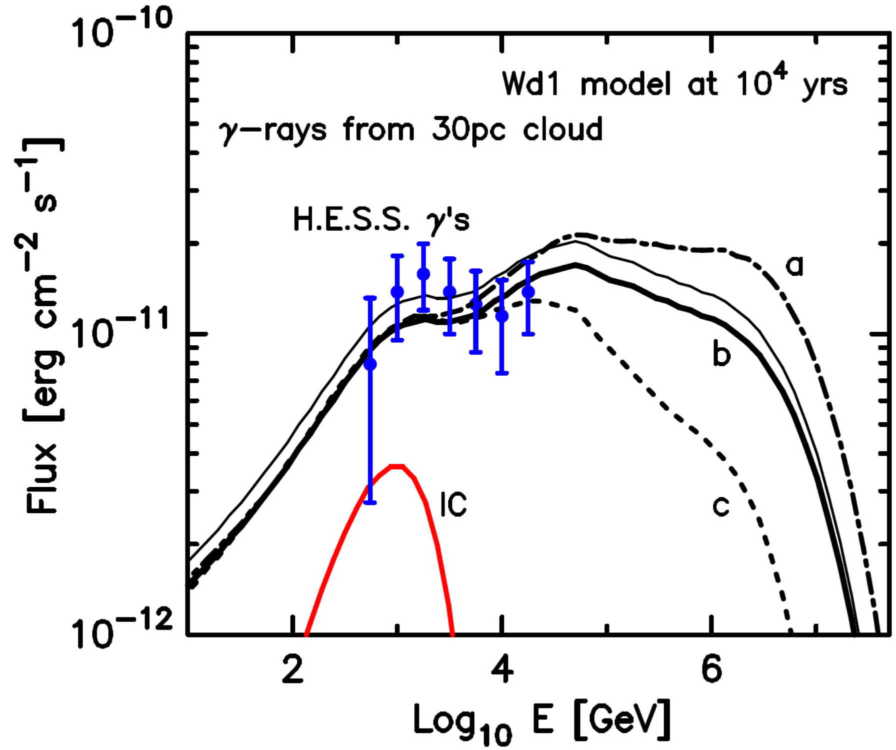
X-ray image of Westerlund 1
(chandra)

- non-thermal Radio and X-rays show hints of particle accelerations

Acceleration site?

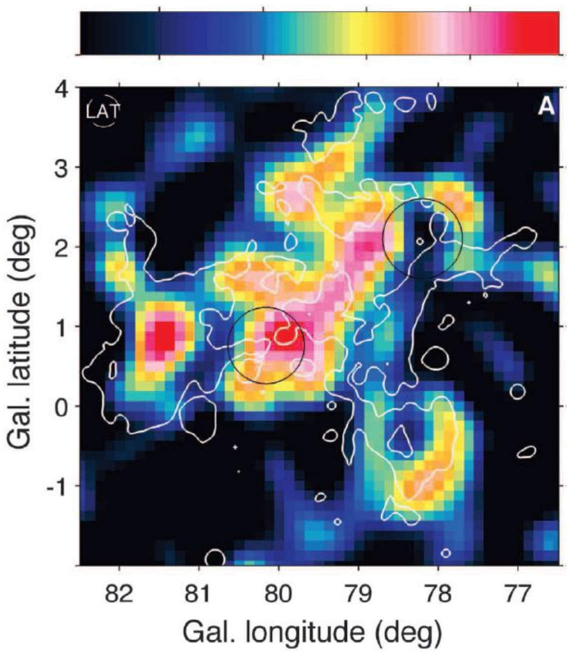


Bykov et.al 2014

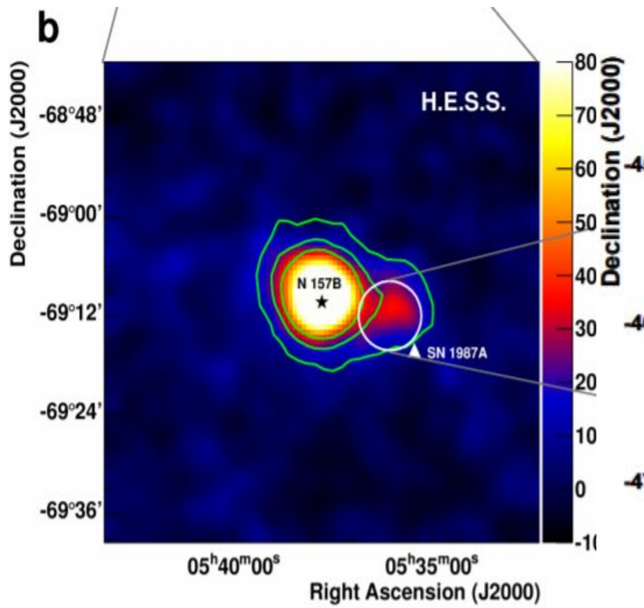


- One promising site: SNR shock colliding with wind termination
- Can accelerate to PeV

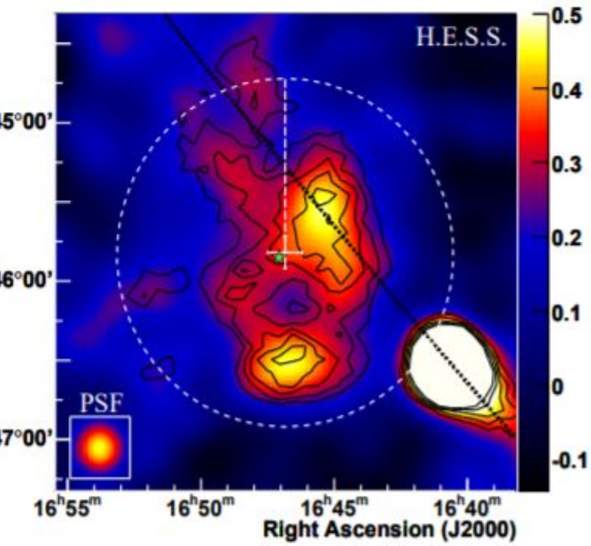
Alternative sources： Young massive clusters



Cygnus Cocoon
(Fermi Collaboration 2012)



30 Doradus C
(H.E.S.S Collaboration 2015)



Westerlund 1
(H.E.S.S Collaboration 2011)

Source population

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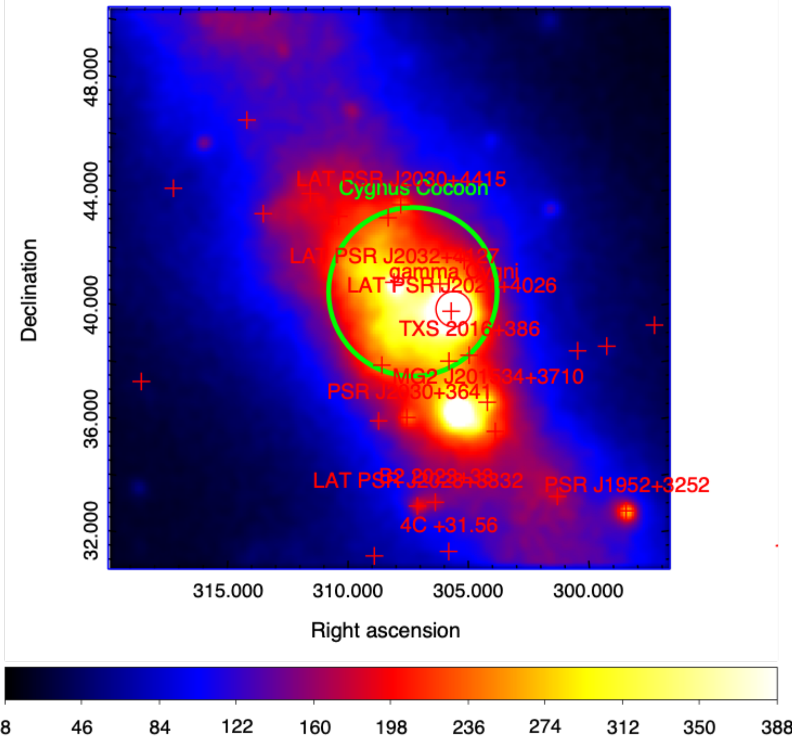
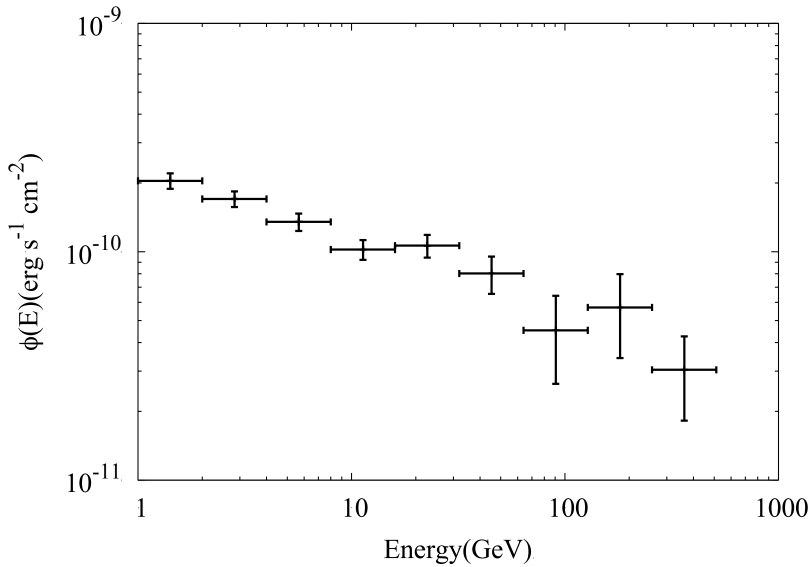
- Cygnus Cocoon



Fermi LAT counts map (>1 GeV)



Spectral Index of -2.2, up to 500 GeV with PASS 8 data.

- More: NGC3603 (Yang & Aharonian 2017), Westerlund 2 (Yang et.al 2018), W43 (Yang & Wang 2020), W40 (Sun et.al 2020) RSGC 1 (Sun et.al 2020)… and more to be discovered and investigated

- All reveal extended gamma-ray emission and hard (2.3 type) gamma-ray spectra

Radial distribution of Cosmic Rays

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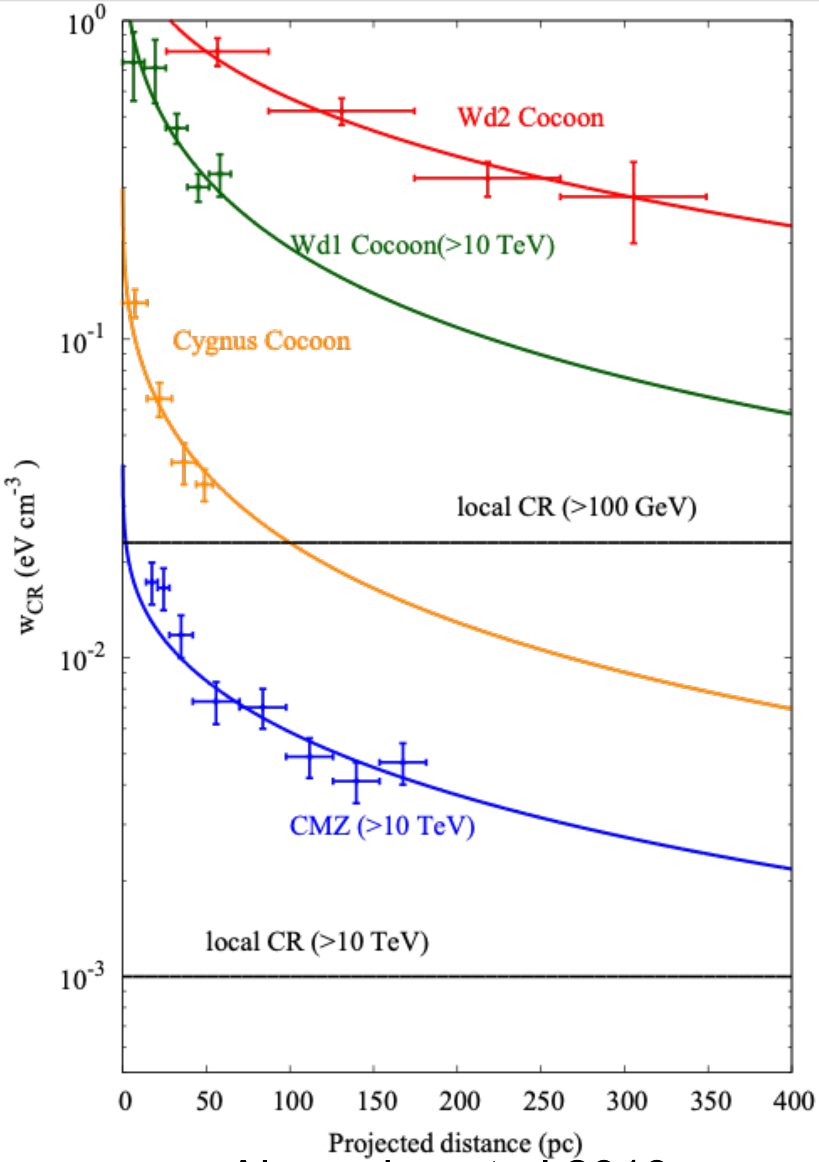

Aharonian et.al 2019

- CR distribution derived by gamma-ray profile and gas distributions
- All four sources (Wd1, Wd2, Cygnus cocoon, GC) show 1/r distribution of CRs
- In diffusion, 1/r profile implies a continuous injection (in the lifetime of clusters)

Massive star clusters

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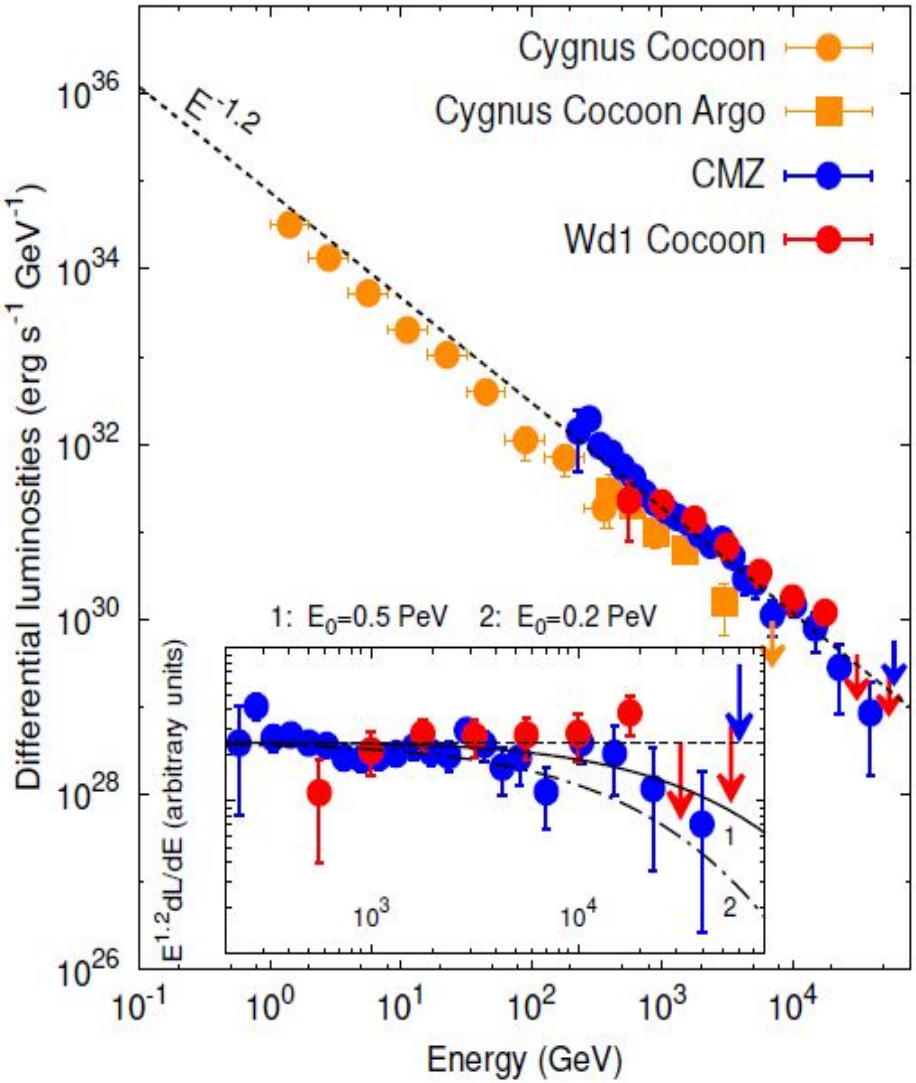

Aharonian et.al 2019

- Cygnus cocoon, Wd 1 and CMZ all emit multi-TeV gamma-ray.
- The spectrum of CMZ and Wd1 put lower limit of cutoff of parent proton spectrum to be several hundred TeV
- PeVatrons?
- LHAASO is the ideal instrument!

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Prospect

PeVatron searching

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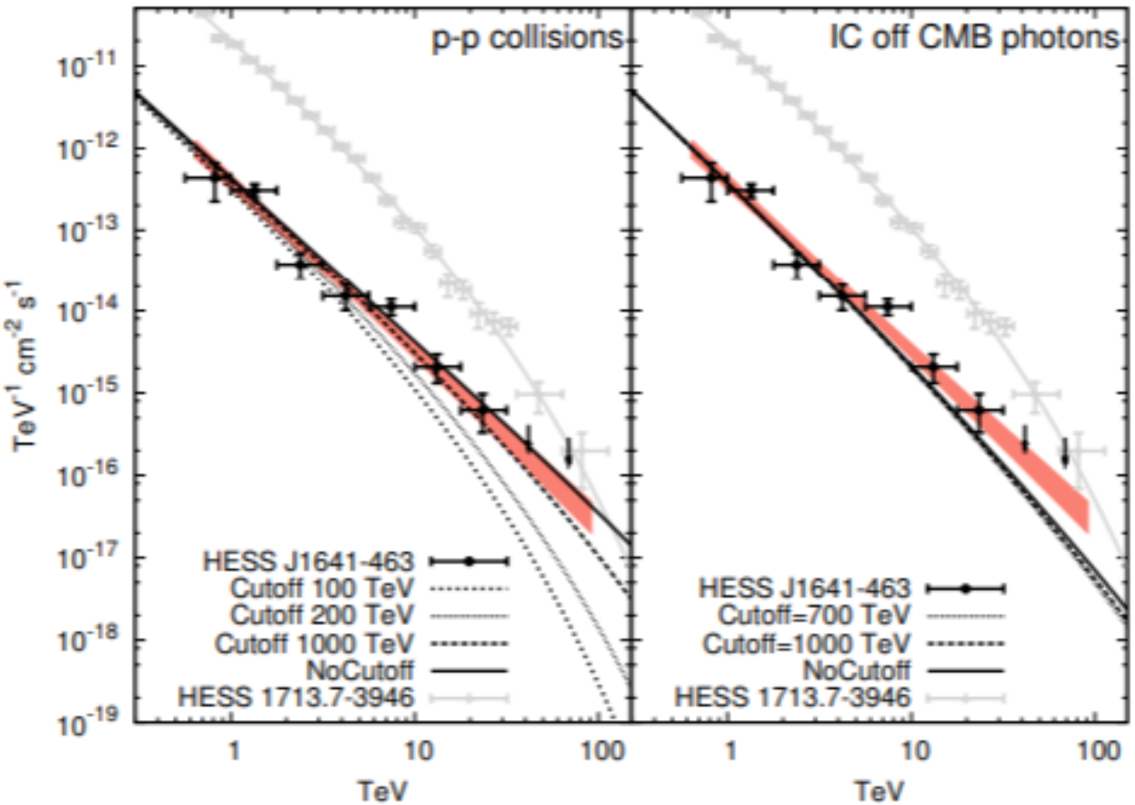
- **Hard gamma-ray spectrum without cutoff can hardly be addressed in leptonic model (cooling and KN effects).**

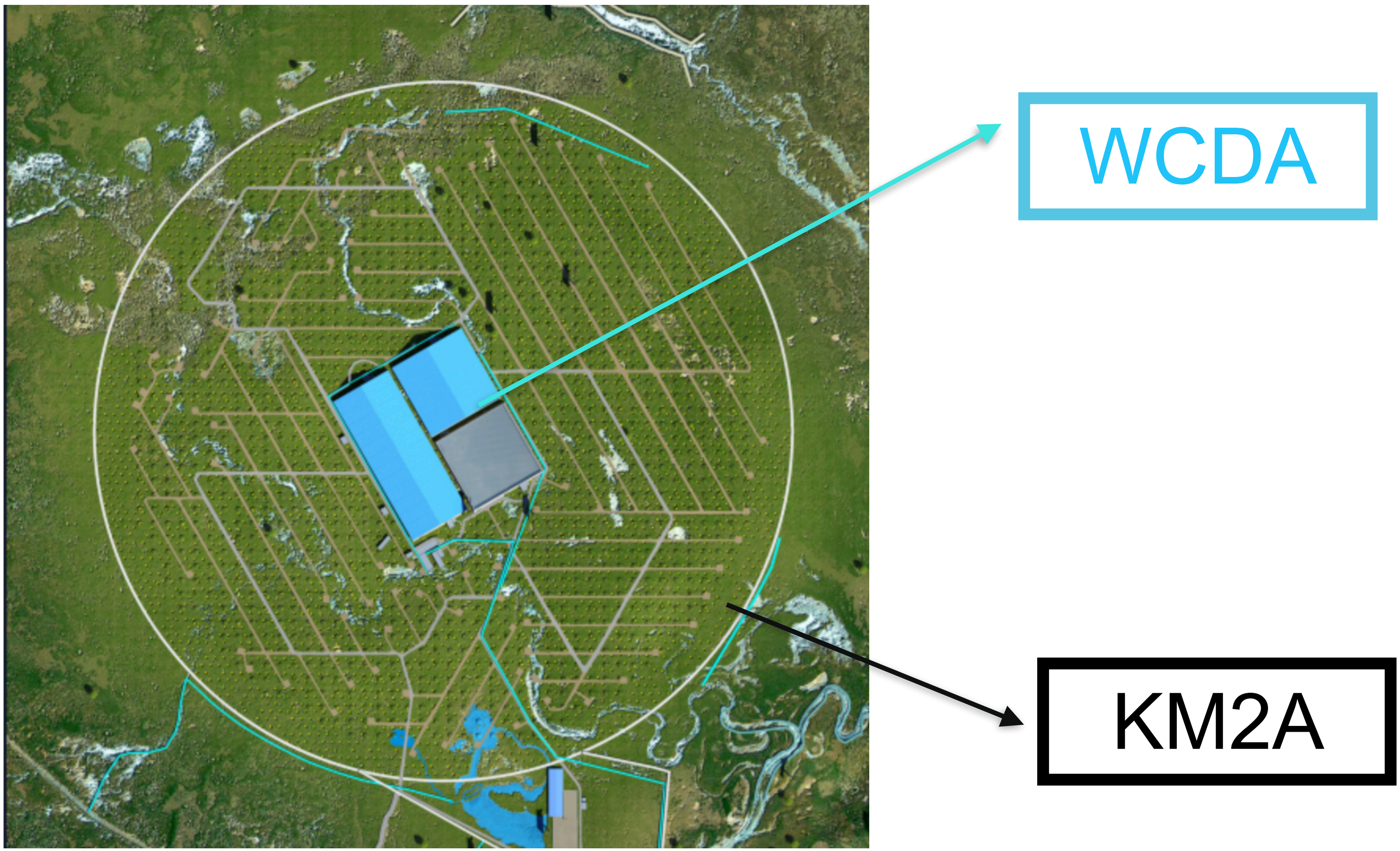
- **no-cutoff in the gamma-ray spectrum up to 25 TeV => no-cutoff in the parent proton spectrum up to ~ PeV.**



Hess J1641-463
(H.E.S.S collaboration 2016)

LHAASO

- Large High Altitude Air ShowerObservatory



LHAASO

- KM2A: 1 km^2 scintillator (ED) and muon detector (MD) array, focus on the ultra high energy gamma-rays (>50 TeV)
- WCDA: Water Cherenkov detector arrays, mainly for TeV gamma-rays
- WFCTA: wide field of view Cherenkov telescope array, measure the shower shape, mainly for the direct measurement of Cosmic rays.

LHAASO sensitivities

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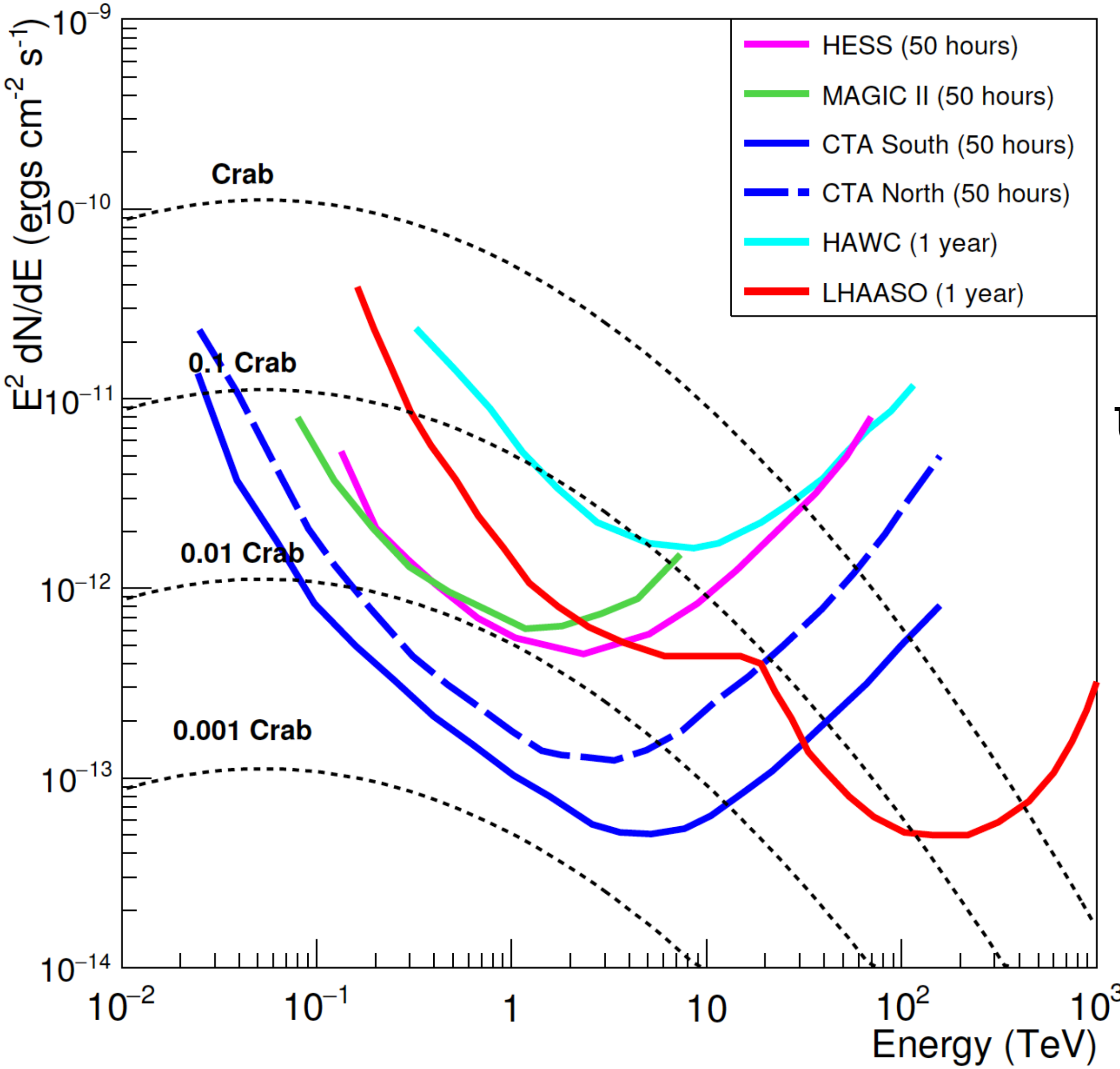

Unprecedented above 50 TeV

ideal PeVatron hunter

LHAASO sensitivity for SNRs

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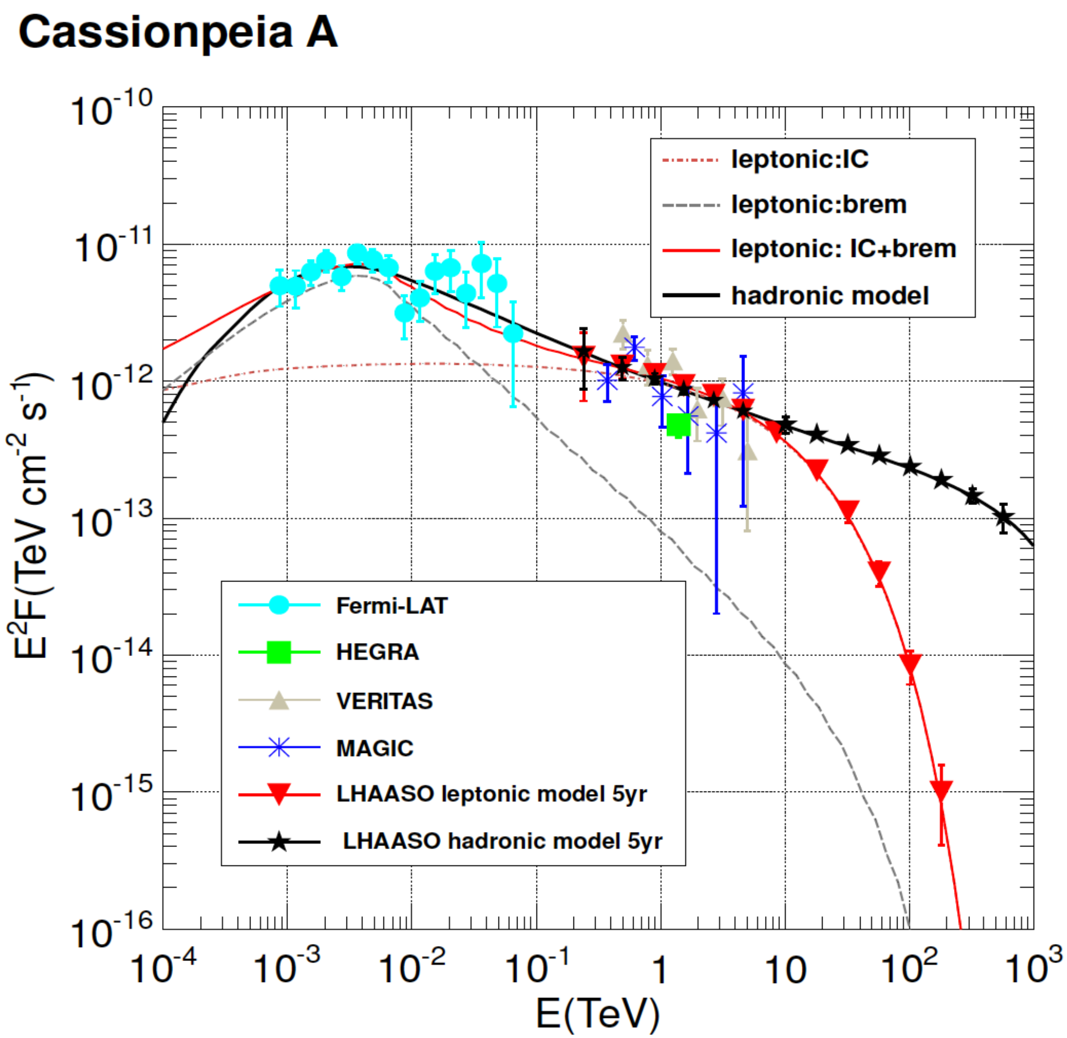

Liu et.al ApJ 826, 1

Concluding Remarks

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- CR distributions from the gamma-ray: origin and propagation
- Low energy end: interplay with the star forming/astro-chemistry, nuclear line may open a new window.
- High energy end : PeVatron searching, LHAASO will play a leading role.

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Thanks!

CCTA Sensitivity

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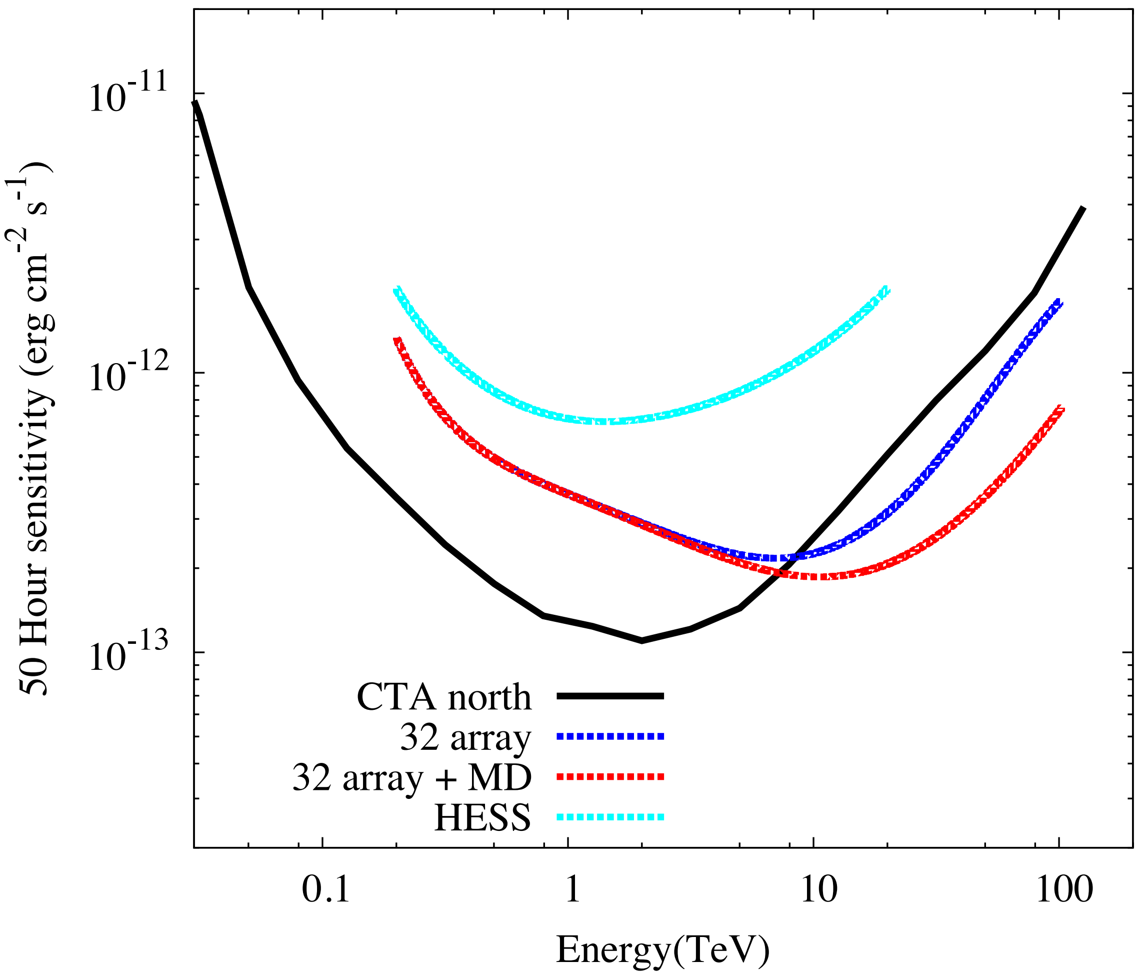
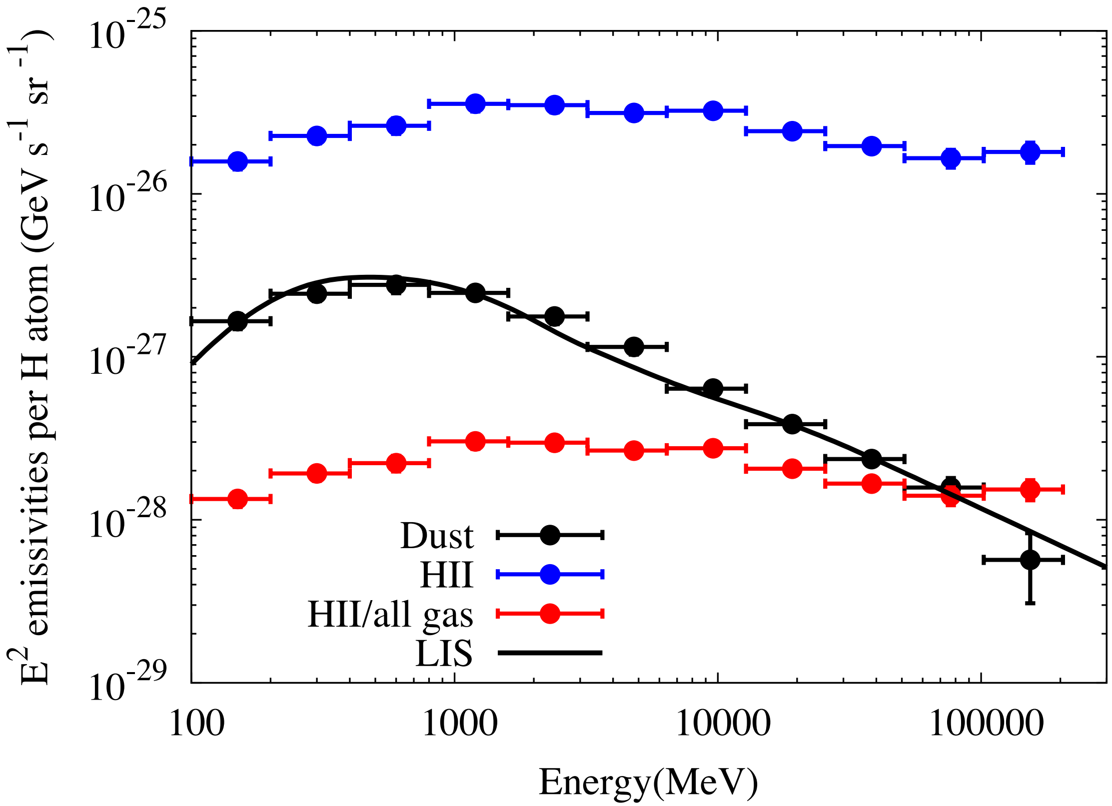


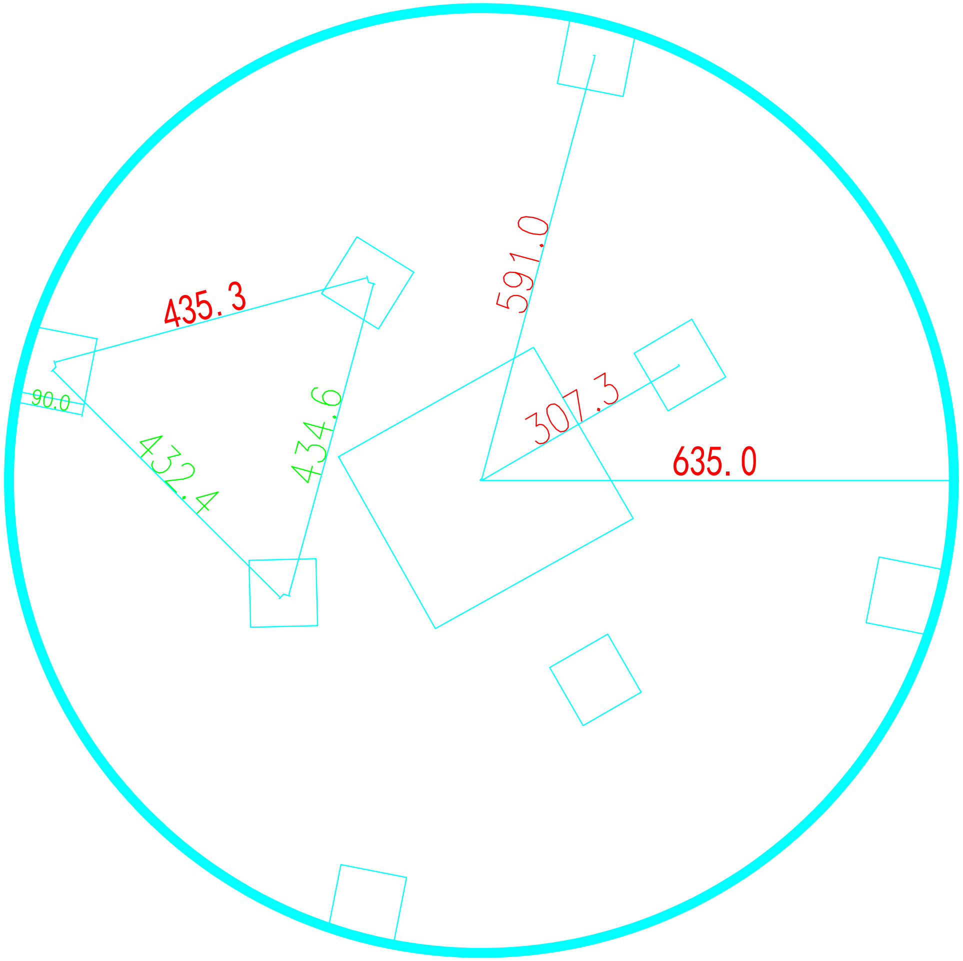
3-times more sensitive than CTA north (under construction) above 10 TeV
order of magnitude better than the running instruments (H.E.S.S, MAGIC, Veritas)

Further test

- use HII gas to trace massive star forming regions (potential CR sources)
- Diffuse gamma can be separated into two components
- One associated with total gas column (dust opacity), with soft spectra. CR “sea”?
- Another with HII gas, with harder spectrum, CR “islands” near sources?



Yang and Liu 2020, in prep



Preliminary plan:

32 Cerenkov telescopes

inside LHAASO site

Further test

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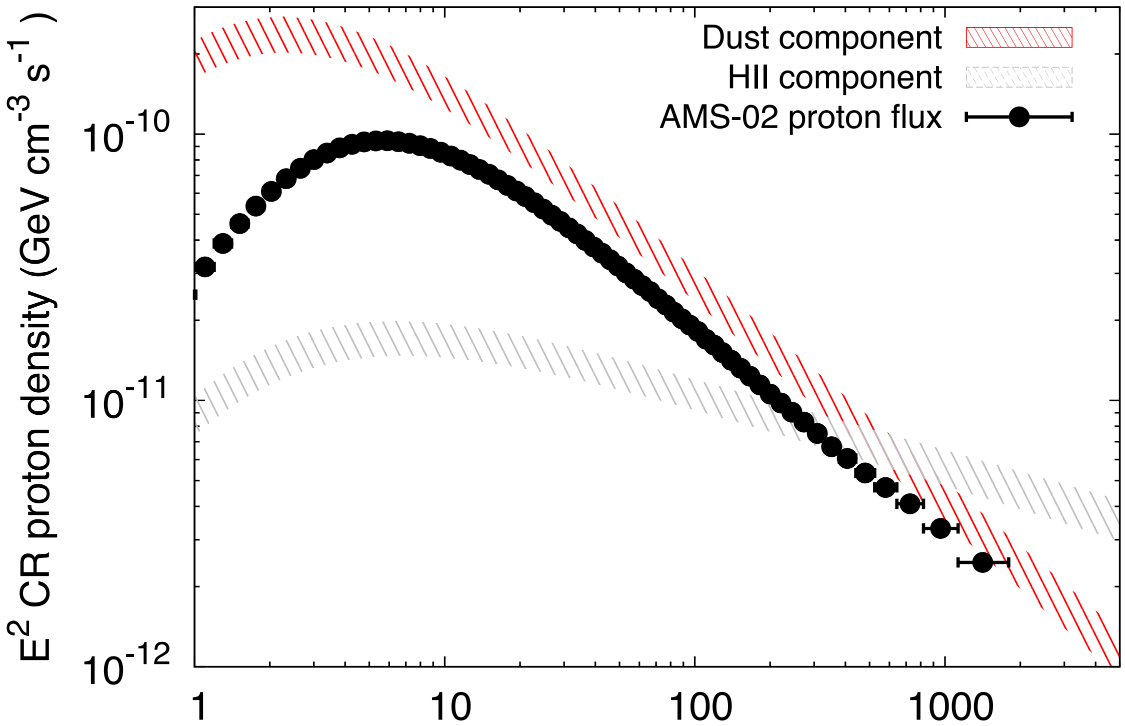

Yang and Liu 2020, in prep